CTRANS
A MONTE CARLO PROGRAM FOR RADIATIVE
TRANSFER IN PLANE PARALLEL ATMOSPHERES
WITH IMBEDDED FINITE CLOUDS — —
DEVELOPMENT, TESTING AND
USER'S GUIDE

Prepared for
National Aeronautics and Space Administration
Goddard Space Flight Center
Greenbelt, Maryland 20771

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SECTION 1.0 INTRODUCTION

This report describes the program called CTRANS which has been designed to perform radiative transfer computations in an atmosphere with horizontal inhomogeneities (clouds). Since the atmosphere-ground system was to be richly detailed, the Monte Carlo method was employed. This means that results are obtained through direct modeling of the physical process of radiative transport. The effects of atmospheric or ground albedo pattern detail are essentially built up from their inpact upon the transport of individual photons. actually tracks the photons backwards through the atmosphere, initiating them at a receiver and following them backwards along their path to the sun as shown schematically in Figure 1.0.1. Backwards tracking has several advantages: the pattern of incident photons generated through backwards tracking automatically reflects the importance to the receiver of each region of the sky. Further, through backwards tracking, the impact of the finite field of view of the receiver and variations in its response over the field of view can be directly simulated. The backwards tracking method, additionally, is well suited for segregating results according to the character of the final scattering suffered prior to entering the receiver.

Section 2.0 describes the backwards tracking Monte Carlo method. In Section 3.0, the results of a range of test applications are reported which serve to illuminate the capabilities and limitations of the program. Section 4.0 contains brief descriptions of each routine comprising CTRANS. Since CTRANS is a highly generalized program with many operational modes and a variety of data structures, Section 5.0 outlines the input parameters. Finally, Section 6.0 contains a FORTRAN listing of CTRANS.

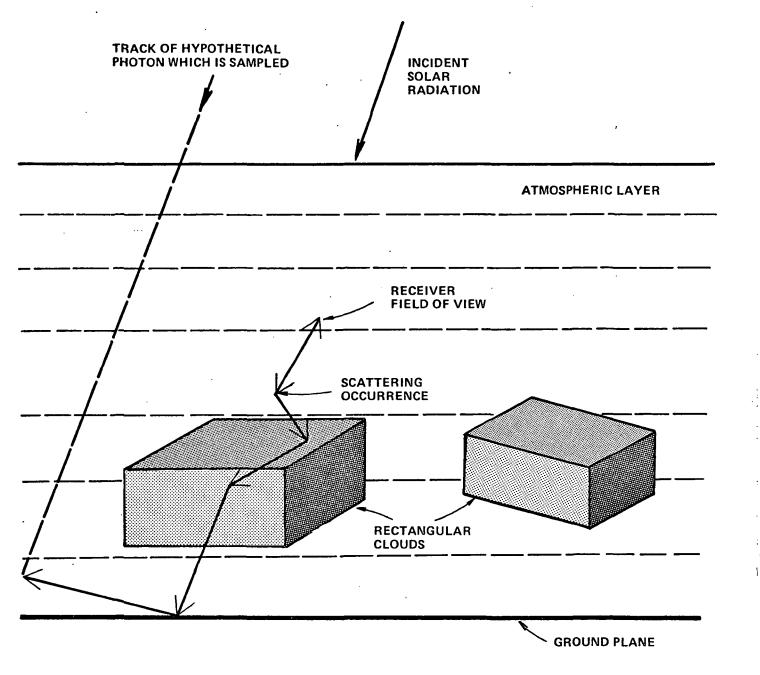


FIGURE 1.0.1. SCHEMATIC REPRESENTATION OF THE PROBLEM GEOMETRY AND THE BACKWARDS MONTE CARLO METHOD

SECTION 2.0 MONTE CARLO METHOD

CTRANS employs the Monte Carlo method to compute flux or intensity upon a receiver situated in or above a plane parallel atmosphere. Photons are tracked backwards through the atmosphere from the receiver along their path to the sun.

The particular advantage offered by the Monte Carlo method lies in the degree of complexity of the earth-atmospheric system which can be modeled.

The present code can handle a plane parallel atmosphere with up to 100 homogeneous layers. Each layer can be composed of up to five distinct scattering (or absorbing) species: molecular scatterers, monodisperse or polydisperse aerosols or hazes, ozone, etc. Further, there may be up to 10 finite rectangular solid clouds which are superimposed upon the layered atmosphere. Each cloud may be composed of any of the five allowed active atmospheric components with an independently specified density for each cloud. Thus, the atmosphere may be inhomogeneous both vertically and horizontally.

The treatment of ground reflection is also generalized. Reflectance at the ground may be either Lambertian or Fresnel in character. Fresnel regions are further characterized by a surface roughness specified by a probabilistic slope distribution (so chosen as to provide an approximate model of a sea surface with waves). The ground plane need not be uniform; rough Fresnel and Lambertian regions may be intermixed in a patchwork fashion as specified by a reflectance type map. Any Lambertian regions may be further broken down into patches having different Lambertian albedos. Results are computed for up to 50 different overlying maps of the Lambertian albedo, simultaneously. Aside from the gain in computational efficiency,

this is important because it provides a means for computing the relative effects of changing the ground albedo pattern with a precision which far surpasses the absolute precision of the Monte Carlo results for a single map.

2.1 PHOTON TRANSPORT AND SAMPLING

CTRANS tracks photons backwards initiating them at the receiver and following them through successive scatterings on their path through the atmosphere. Backwards tracking has been adopted for a number of reasons. It permits the efficient and direct simulation of receivers with finite fields of view and/or finite extent. A forward tracking scheme would require the distribution of initial photons over a wide area when horizontal inhomogeneities are present and this would introduce large variances in the results. By initiating photons at the receiver, this source of variance is largely avoided.

To understand backward tracking, consider first a single scattering event. If $\underline{\underline{I}}_0$ is the Stokes vector of incident unscattered light, $\underline{\underline{S}}$ is the scattering phase matrix and $\underline{\underline{R}}$ is a rotation matrix which serves to rotate the plane of reference of the Stokes vector into the scattering plane, the Stokes vector after scattering $\underline{\underline{I}}_f$ is given by

$$\underline{I}_f = R_r \underline{S} \underline{R} \underline{I}_o$$

where R_r rotates the final Stokes vector into the reference plane of the receiver. After two scatterings, the final Stokes vector will be

$$\underline{I}_f = R_r \underline{S}_2 \underline{R}_2 \underline{S}_1 \underline{R}_1 \underline{I}_o$$

Cast in this way, it is apparent that what is important is the product matrix which is built from the scattering matrices appropriate to the string of scatterings suffered by the photon. This product can be built up either from the left or from the right with the same result; i.e., we may equally well follow the photon's history forward or backwards if we accumulate scatterings properly into the product [which we call the cumulative scattering matrix].

There is a further advantage in tracking backwards: contributions from several solar directions may be accumulated from a single track. For example, if the photon is tracked backwards from the receiver, only the last rotation and scattering contain any reference to solar direction. Replacing one sun by another means changing only these two matrices. For each solar direction, at most, two 4 x 4 matrix multiplications must be performed.

Photons are initialized at the reciever; which may be general in the following ways: it may have finite or infinitessimal (conical) field of view; it may be oriented pointing toward any zenith angle with any azimuthal angle; it may be at any altitude and have any specified x-y coordinate. On the other hand, for flux calculations the receiver aperture may be finite. This is accomplished by letting the receiver be an entire cloud face (any face of any cloud, pointing into or out of the cloud).

First, initial photon coordinates are created. If the receiver is a cloud face, initial photon coordinates are chosen from a uniform distribution over the cloud face. The photon direction is then determined using: 1) uniform distribution over solid angle for a cloud-face receiver, or 2) cosine distribution over the conical field of view of an infinitessimal receiver, or 3) in a specified direction if the field of view

is infinitessimal. For each type of distribution, the appropriate weight is computed (to be multiplied into the sampled values upon sampling).

The distance to be traversed to a scattering is computed on the basis of optical distance to be traversed: if ρ is a uniformly distributed random number, the optical distance to be traversed is given by

$$\tau = - \ln \rho \tag{2.1.1}$$

Through reference to a table containing the optical thickness of each layer of the ambient atmospher and taking into account the added optical density in the interior of clouds, the optical distance traversed and physical distance traversed are incremented until an optical distance τ has been reached. If, before this has occured, the photon exits the atmosphere, the photon is terminated. If the photon strikes the ground, the scattering will be a ground scattering and a ground scattering indicator is set (ISCAT=1).

Having located the position of a scattering, the scatterer type is chosen on the basis of the relative contribution of each scattering species at that point to the local optical density. Here, a cumulative distribution function, which has been tabulated for each layer of the ambient atmosphere, is used in conjunction with any additional optical density due to the presence of a cloud. Photons are not explicitly absorbed. Instead, a weight is accumulated which at any point along the path reflects the probability that the photon has not been absorbed.

After having determined the scatterer type, the probability that the photon could have come from each sun is sampled. If $\underline{\underline{S}}_{I}$ is the cumulative scattering matrix at the point of sampling, the sampled Stokes vector will be

$$\underline{\underline{I}}_{S} = \underline{\underline{S}}_{\underline{\underline{I}}} \underline{\underline{S}}_{\underline{\underline{J}}} \underline{\underline{J}}_{\underline{\underline{O}}} e^{-t} \underline{j} w \qquad (2.1.2)$$

Where R_j is a rotation matrix appropriate to a scattering into the j^{th} solar direction; S_j is a scattering matrix (a Mueller matrix if the scatterer is an atmospheric constituent); I_0 is the incident Stokes vector; t_j is the total optical distance along the path connecting the scattering point and \sup_j ; w is a weight reflecting the probability that the photon has not been absorbed (including ground absorption for Lambertian regions encountered). There will be a different value for w for each Lambertian albedo map. Note that the above form assumes that \underline{I}_0 represents intially unpolarized light. If the incident light from the sun were polarized, an additional rotation would be necessary. R_j has the form

$$R_{j} = \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & \cos 2X & \sin 2X & 0 \\ 0 & -\sin 2X & \cos 2X & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$$
 (2.1.3)

where X is the rotation angle between the normal to the previous scattering plane (or the reference direction in the case of first scatterings) and the normal to the scattering plane for a scattering into the jth solar direction. S_j has the general form

$$S_{j} = \frac{1}{4\pi} \begin{pmatrix} S_{11} & S_{12} & 0 & 0 \\ S_{12} & S_{22} & 0 & 0 \\ 0 & 0 & S_{33} & S_{34} \\ 0 & 0 & -S_{34} & S_{44} \end{pmatrix}$$
(2.1.4)

where $S_{\ell,m}$ depends upon scattering angle. For Rayleigh scattering, $S_{11} = S_{22} = \frac{3}{4} (1 + \cos^2 \beta_i)$

$$S_{12} = \frac{3}{4} (\cos^2 \beta_1 - 1)$$

$$S_{33} = S_{44} = \frac{3}{2} \cos \beta_{j}$$

$$S_{34} = 0$$

where β_j is the scattering angle for a scattering into the j^{th} solar direction. For Lambertian scattering at the ground, $S_{11} = 4\cos\theta_j$, where θ_j is the zenith angle of the j^{th} sun and all other matrix elements are zero. Lambertian reflection thus is perfectly depolarizing.

Mie scattering and Fresnel scattering both have matrices with the form of Eq. (2.1.4). Mie elements are stored in tabular form and values for particular scattering angles are obtained by interpolation. The Fresnel matrix elements are given in the discussion of ground reflectance.

The next step after sampling is to accumulate a "real" scattering. Knowing the scatterer type, an appropriate scattering angle is found. If the scattering angle were chosen from a uniform distribution, and if the scattering matrix appropriate to the medium exhibited sharp peaks, then most photons would rapidly lose weight and large numbers of photon tracks would be necessary in order to achieve a reasonably small variance. To resolve this difficulty, we use the concept of importance sampling. We choose the scattering direction from a biased distribution designed to maximize the resultant contribution. The bias thus introduced is then removed by applying a weight. More concretely, we wish to fold in a scattering matrix

$$S = \begin{pmatrix} S_{11} & S_{12} & 0 & 0 \\ S_{12} & S_{22} & 0 & 0 \\ 0 & 0 & S_{33} & S_{34} \\ 0 & 0 & -S_{34} & S_{44} \end{pmatrix} \frac{dw}{4\pi}$$

$$= \frac{1}{f(\beta)} \begin{pmatrix} S_{11} & S_{12} & 0 & 0 \\ S_{12} & S_{22} & 0 & 0 \\ 0 & 0 & S_{33} & S_{34} \\ 0 & 0 & -S_{34} & S_{44} \end{pmatrix} \frac{f(\beta) dw}{4\pi}$$
(2.1.5)

Instead of choosing the scattering angle from a probability distribution corresponding to the probability density $\frac{dw}{4\pi}$ (uniform), we use the distribution corresponding to the probability density $\frac{f(\beta)}{4\pi}$. The weight of the photon following the scattering can thus be maximized by choosing $f(\beta) = S_{11}$ times a constant chosen so that the first matrix element is always unity. A real scattering is accumulated by multiplying into the cumulative scattering matrix first a rotation, and then the scattering matrix.

The process is repeated until either a designated maximum number of scatterings have been suffered, the photon leaves the atmosphere or its weight drops below the assigned threshold (due, for example, to atmospheric absorption).

2.2 ENCOUNTERS WITH CLOUDS

The presence of finite clouds complicates the evaluation of the fate of the path of the photon in that there is no simple way to determine if a given straight line in space enters a specified rectangular box. It was therefore necessary to undertake an analysis to find optimum procedures for evaluating the path as quickly as possible.

This analysis is applied in two areas of the program:

- 1. In evaluating the path of the "backward" photon to the "previous" scattering.
- 2. In computing the probability that a "forward" hypothetical photon would reach a given scattering point from the Sun without scattering.

The problem has been treated in two parts. First, we have found a quick and efficient method for eliminating clouds which could not possibly be entered by the photon path, and secondly, we have performed a detailed analysis to evaluate the uneliminated clouds.

There is a simple method for determining whether it is possible for a straight line to enter a given rectangular cloud. The approach is to circumscribe the cloud with a sphere, the center of which coincides with the center of mass of the cloud. The radius, r, of the sphere can be taken to be $(a^2+b^2+c^2)^{1/2}$ where the lengths of the sides of the cloud are (2a,2b,2c). In this case the eight corners of the box lie on the sphere as illustrated in Figure 2.2.1.

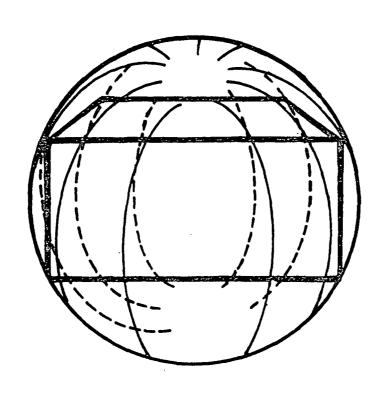


FIGURE 2.2.1. ILLUSTRATION OF CIRCUMSCRIBING SPHERE CONCEPT

Since the sphere completely contains the cloud then if a line cannot enter the sphere, it cannot enter the cloud. Thus, the first check should be to find if the path can enter any cloud circumscribing spheres. We shall present the analysis for a single cloud with the understanding that this must be repeated for all clouds in the program.

Suppose we have a line with direction cosines (c_x, c_y, c_z) leaving a point $A(x_1, y_1, z_1)$. The equation of the line is

$$\frac{x-x_1}{c_x} = \frac{y-y_1}{c_y} = \frac{z-z_1}{c_z}$$
 (2.2.1)

The distance, R, of the point O(X,Y,Z) from A is

$$R = \{(X-x_1)^2 + (Y-y_1)^2 + (Z-z_1)^2\}^{1/2}$$
(2.2.2)

and the point of closest approach of the line to 0 is such that AP is perpendicular to OP, as illustrated in Figure 2.2.2. Now the cosine of the angle OAP is given by the dot product of the direction cosines of AP and AO; i.e.,

$$\cos OAP = \frac{(X-x_1)c_X}{R} + \frac{(Y-y_1)c_Y}{R} + \frac{(Z-z_1)c_Z}{R}$$
 (2.2.3)

But the distance AP is given by R cos OAP; i.e.,

$$AP = (X-x_1)c_X + (Y-y_1)c_Y + (Z-z_1)c_Z \qquad (2.2.4)$$

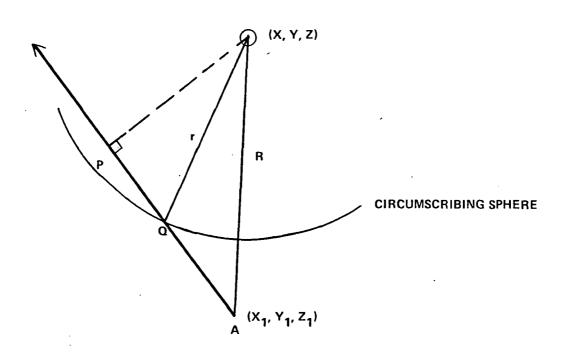


FIGURE 2.2.2. ILLUSTRATION OF POINT OF CLOSEST APPROACH OF A LINE TO A POINT

and by Pythagorus's theorem,

$$OP^{2} = R^{2} - AP^{2}$$

$$= R^{2} - \{(X-x_{1})c_{x} + (Y-y_{1})c_{y} + (Z-z_{1})c_{z}\}^{2}$$
(2.2.5)

Now, if OP^2 is greater than r^2 , the squared radius of the circumscribing sphere, then the path cannot enter the sphere. Also, if the quantity in curly parentheses is negative, then again entry is impossible.

Note that ${\sf OP}^2$ can be computed in the general coordinate system and that no square roots or complicated functions are necessary for its evaluation.

If we are tracking the path of the backward photon then another check can be made. Although the point of closest approach may be inside the circumscribing sphere, it may be that the length, L, of the path is insufficient to reach the sphere. Suppose, in Figure 2.2.2 that the line reaches the sphere at Q. Then OQ has length r, the radius of the sphere. Now from (2.2.4) and (2.2.5) we are able to evaluate both OP and AP and by Pythagorus's theorem we have that

$$QP^2 = r^2 - OP^2 (2.2.6)$$

Thus we are able to evaluate the distance to the sphere as

$$AQ = AP - QP \qquad (2.2.7)$$

which may be evaluated by (2.2.4) and (2.2.6). Testing this value against L gives us another possible means of cloud elimination.

We have attempted to use the circumscribing sphere concept further, but we finally concluded that there are some dangers in its further application. For example, the sequence of distances along the path to a cloud need not be the same as the sequence of distances to the circumscribing spheres or the point of closest approach. An illustration of such a case is given in Figure 2.2.3. The reason for the difficulty is that the circumscribing sphere gives no information as to cloud orientation and can therefore mask orientation differences which may be important. For this reason we proceed in the program at this point to a completely general analysis.

The first step in treating an uneliminated cloud is to convert to the coordinate system of the cloud. Thus the location of the scattering point (x_1,y_1,z_1) is transformed to (x_{1c},y_{1c},z_{1c}) using

$$x_{1c} = (x_{1}-X) \cos \theta_{c} + (y_{1}-Y) \sin \theta_{c}$$
 $y_{1c} = -(x_{1}-X) \sin \theta_{c} + (y_{1}-Y) \cos \theta_{c}$
 $z_{1c} = z_{1}-z$

(2.2.8)

where θ_{C} is the orientation angle of the cloud and (X,Y,Z) are the coordinates of the center of mass of the cloud. The direction cosines of the line are transformed to the cloud system using

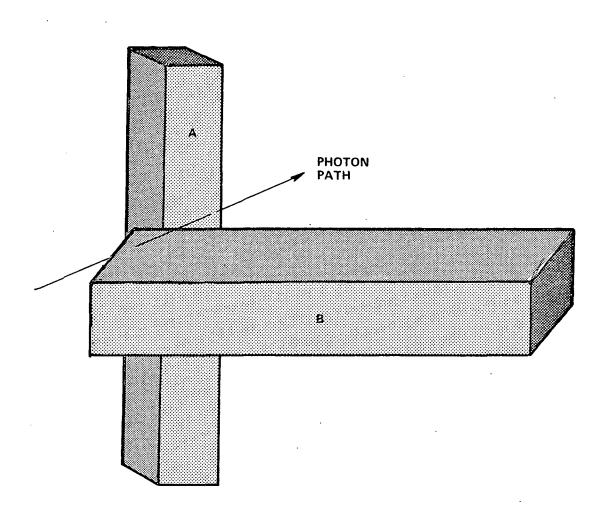


FIGURE 2.2.3. EXAMPLE OF SEQUENCING FAILURE OF CIRCUMSCRIBING SPHERES ALGORITHM

$$C_{cx} = C_{x} \cos \theta_{c} + C_{y} \sin \theta_{c}$$

$$C_{cy} = -C_{x} \sin \theta_{c} - C_{y} \cos \theta_{c}$$

$$C_{cz} = C_{z}$$

$$(2.2.9)$$

We shall present the analysis for determining if the photon enters a side perpendicular to the cloud x axis. In the computer program we index terms relating to the cloud x, y and z axes 1, 2 and 3, respectively. The analysis for the sides perpendicular to the y and the z axes can then be performed by merely rotating the ordered index set (1,2,3) to (2,3,1) and (3,1,2), respectively.

The following tests will be performed in sequence.

1. Is $|x_{1c}| \le a$, the half length of the edges parallel to the cloud x axis. (2.2.10)

If this condition is satisfied, then the photon cannot possibly enter a face perpendicular to the cloud x axis.

2. If $x_{1c} < -a$, then the path can enter the face $x_{c} = -a$. In this case, set q = -a.

If $x_{1c} > a$, then the path can enter the face, $x_c=a$. In this case, set q=a.

3. Compute $(x_{1c}-q) \cdot C_{cx}$. If this is positive, then the photon is traveling away from the face, $x_c=q$, and will not enter the cloud.

4. Compute the (y_c, z_c) values of the point of intersection of the line with the plane, $x_c=q$. These are given by

$$y_{cI} = y_{1c} + C_{cy} \frac{q - x_{1c}}{C_{cx}}$$

$$z_{cI} = z_{1c} + C_{cz} \frac{q - x_{1c}}{C_{cx}}$$
(2.2.11)

5. Check to see if both the conditions (2.2.12) apply.

$$|y_{cI}| < b$$
 (2.2.12) $|z_{cI}| < c$

If these are both true, then the photon enters the cloud at the point (q, y_{cI}, z_{cI}) . Otherwise, entry does not occur at the x=q plane.

When it is determined that a photon has entered a given cloud then the tracking mechanism from that point on will be performed in the cloud coordinate system.

If it is determined that a later scattering point (x_{1c}, y_{1c}, z_{1c}) lies outside the cloud then we determine the exit point using steps (4) and (5) only.

For evaluating the track of the "forward" hypothetical photon from the Sun to a scattering center we will determine if the direction cosine $C_{\rm cxs}$ of the line from the scattering point to the Sun is positive or negative. (The solar direction cosines ($C_{\rm cxs}$, $C_{\rm cys}$, $C_{\rm czs}$) will be computed for each cloud prior to the simulation.) If $C_{\rm cxs}$ is positive, we set q=a, whereas if it is negative, we shall set q=-a. At this point steps (4) and (4) are employed to find the point of exit from the cloud.

2.3 GROUND SCATTERING

CTRANS has the capability to model reflections from a spatially inhomogeneous ground. In the present version of the code, the physical mechanism of reflection may be modeled as either Lambertian reflection or as Fresnel reflection from a rough surface and these modes may be intermixed in a patchwork fashion.

Reflectance type is basically controlled by a reflectance type map which dictates the reflectance character from point to point across the infinite ground plane. When a photon strikes the ground, its impact coordinates are computed and then used to decide whether the point of impact was within a Lambertian or a rough Fresnel region.

Since Lambertian reflection is not physically dependent upon the direction of incidence or reflection but only upon a scalar Lambertian albedo, CTRANS can model simultaneously the effect of several values (up to 50) of the albedo for each of the Lambertian regions. This is accomplished in the following way: A vector of scalar weights is maintained; one for each albedo map. When a Lambertian scattering occurs, the albedo at the point of impact is determined for each of the albedo maps and this is multiplied into the corresponding albedo weight to obtain a new weight for subsequent use. Upon sampling, each current albedo weight is folded into the sampled Stokes vector to obtain samples corresponding to each of the allowed albedo maps. If the impact point lies within a Fresnel region, however, the vector of albedo weights is left unchanged (since it represents the cumulative effects of only the Lambertian regions).

Symbolically, the entire reflection process is as follows: given an old cumulative scattering matrix \mathbf{S}_{o} , the new one, \mathbf{S}_{N} , is computed as:

$$S_{N} = S_{O} R s$$

where R is a rotation matrix and s depends upon the reflectance type at the impact point:

where S_{ij} represent Fresnel matrix elements:

$$S_{11} = (R_{11}^2 + R_1^2)/2$$

 $S_{12} = (R_{11}^2 - R_1^2)/2$
 $S_{33} = R_{11} R_1$

and

$$R_{11} = \frac{n \cos \chi_i - \cos \chi_t}{n \cos \chi_i + \cos \chi_t}$$

$$R_1 = \frac{\cos \chi_i - n \cos \chi_t}{\cos \chi_i + n \cos \chi_t}$$

n = relative index of refraction of the surface

 χ_i = angle of incidence

 χ_t = angle of refraction, computed from χ_i through Snell's law at the interface: n sin χ_t = sin χ_i .

The vector of scalar weights is renewed according to

$$A_i^{\text{new}} = A_i^{\text{old}} \times a_i(J;X,Y)$$

X,Y = coordinates of impact point

J=0 = region is Lambertian and the values of $a_{i}(i=1,50)$ are determined from the Lambertian albedo maps according to the coordinates (X,Y)

 $J=1 = region is Fresnel a_i = 1(i=1,50).$

The two kinds of ground maps (Lambertian albedo maps and the reflectance type map) are specified in CTRANS in two different ways.

The Lambertian albedo maps are specified in a fixed form by the subroutine ALBEDO. To change albedo map structures one version of ALBEDO is interchanged with another. The possible variety of maps is thus limited only by the imagination. ALBEDO, in addition to its function of providing specific values of the scalar weights at a specified point has the function of writing out text to describe the various maps.

The reflectance type map is specified by options read in by subroutine INPUT and is therefore less flexible than the albedo maps. As options, the following choices are presently available:

- 1) Uniform plane.
- 2) Adjoining half planes of different reflectance type bounded by a line parallel to the y-axis whose position is specifiable by its x-coordinate.
- Non-overlapping rectangles set upon a uniform background (up to 20 such rectangles are allowed).

In all of the above, the background is specifiable as either Lambertian or rough Fresnel with inhomogeneities having either rough Fresnel or Lambertian character, respectively. The rectangles are specified by the (X,Y) coordinates of their centers and their X-Y edge lengths.

Text describing the chosen pattern is written out by INPUT. A cautionary note is in order here. Since the two types of maps are specified by unrelated means, care must be taken in the interpretation of results: The Albedo maps only apply to regions specified to have Lambertian character. Thus, using 50 varigated albedo maps and specifying the reflectance type map as a uniform rough Fresnel plane would produce identical results for all 50 albedo maps.

For rough Fresnel regions, the roughness is characterized as a distribution of randomly oriented sloping plane facets. The probability density for the distribution of slopes was chosen to have two-dimensional Gaussian form. $p(Z_x, Z_y) \times \delta Z_x \delta Z_y$ is the fraction of a small horizontal unit area of surface for which the x,y components of the slope are within the limits $Z_x + \frac{1}{2} \delta Z_x$ and $Z_y + \frac{1}{2} \delta Z_y$. We specify $p(Z_x, Z_y)$ as:

$$p(Z_x, Z_y) = (\pi\sigma^2)^{-1} \exp[-(Z_x^2 + Z_y^2) / \sigma^2]$$

The variance, σ^2 is specified either by three parameters W0, W1, W; or as a set value specified on input. When specified through parameters, σ^2 is given as

$$\sigma^2 = W0 + W1 \times W$$

W represents wind speed (m/sec) in a model for the surface of the sea $^{1)}$ and W0, W1 have either their default values (W0 = .0015, W1 = 2.54 x 10^{-3}) or values specified by card input. In the present code, all rough Fresnel regions share the same parametric values. This restriction could be relatively easily relaxed if desired.

The above form for $p(Z_x, Z_y)$ represents that applicable for a photon incident from the vertical. For photons incident from some other direction, the distribution must be modified. Let the photon's direction be given by unit vector \hat{k} , the direction of the normal to a facet be \hat{n} , and the vertical direction be \hat{Z} . The probability of encounter will then be given by:

$$\frac{\hat{k}\cdot\hat{n}}{(\hat{n}\cdot\hat{z})(\hat{k}\cdot\hat{z})} p(\hat{n}) .$$

Having determined a slope at the photon impact point, the angle of incidence and, thence, the Fresnel matrix elements can be determined. If k and k are the directions of propagation of the incident and reflected photons, their relation to n is

$$\frac{\hat{k}' - \hat{k}}{|\hat{k}' - \hat{k}|} = \hat{n} .$$

When sampling the contributions of unscattered photons from the Sun at a Fresnel scattering point, the direction of reflection is given by the solar direction and one merely evaluates the probability of having encountered a slope having the proper orientation. Note that in the present model the location of the Sun is given by a Dirac delta function (incoming plane waves) rather than being spread over a small angular region. The implied integration over the solar direction (or, equivalently, the change of variables within the delta function specifying the solar direction) introduces an additional factor of

.25
$$(\hat{n} \cdot \hat{z})^{-3} (k \cdot \hat{n})^{-1}$$

Our present formulation does not properly account for shadowing (waves whose surfaces are inaccessible due to their orientation) or for hiding (slopes inaccessible because of intervening higher waves). Thus, we expect some systematic bias, especially when either the sun or the receiver is low in the sky.

SECTION 3.0

TESTING: CAPABILITIES AND LIMITATIONS
OF CTRANS

CTRANS has been tested and validated through comparison with the results of other calculations under a variety of conditions. Here, the results of representative cases will be presented and discussed.

The basic functional structure of the Monte Carlo code (backwards-tracking algorithm, Lambertian ground reflection, statistical normalizations, etc.) may be tested via calculations of radiative transfer through a pure Rayleigh atmosphere and comparison with the well established results of Coulson, Dave, and Sekera. Next, the atmosphere is made more complicated by adding layered structure and additional active atmospheric species. This provides a basic test of the algorithms for scatterer selection, the treatment of absorption, and the treatment of polydispersions. There may be special problems associated with optically thick atmospheres composed of strongly anisotropic scatterers so this case is of interest.

CTRANS' capacity to handle horizontal inhomogeneities is basically evaluated by means of a calculation with a single cloud in the atmosphere and comparison with the results of McKee and Cox. 3) In this test, some difficulties were detected which are associated with the calculation of flux (in optically thick clouds) by means of a backwards tracking code. These difficulties are partially resolved by using a specially adopted forwards-tracking code which will be briefly described.

CTRANS' capacities to treat horizontal ground reflection inhomogeneities are presented through test results. Absolute validation is difficult here in the absence of dependable analytical results, but having validated the microscopic reflection mechanism (through the inspection of results for uniform ground planes) only problems associated with the statistical density of sampling might be present. These statistical problems may be detected through the comparison of results from statistically independent runs.

3.1 HORIZONTALLY HOMOGENEOUS ATMOSPHERES

The first and most basic test case is that for a horizontally (and vertically) homogeneous Rayleigh atmosphere. Here CTRANS results may be compared with those of Coulson, Dave and Sekera ²⁾ who have tabulated values for the Stokes parameters and polarization for Rayleigh atmospheres with optical thicknesses ranging up to 1.0 and for uniform Lambertian ground planes with albedos from 0.0 to 0.8.

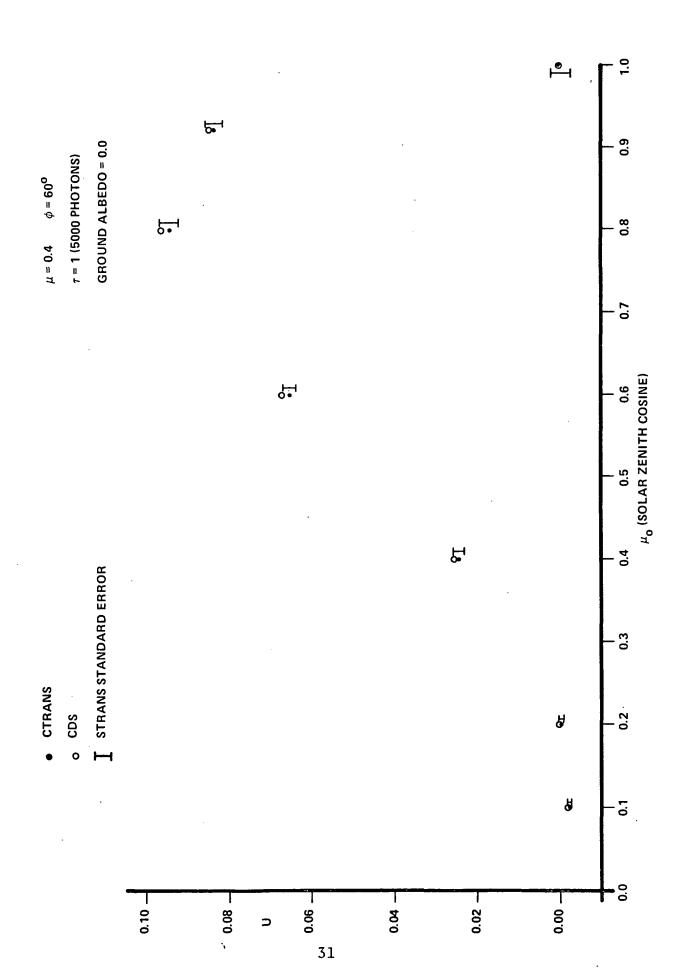
We have run tests over a range of optical thicknesses. Of these, the most stringent test is provided by the largest optical thickness (τ = 1). Since, in this case, single scattering, multiple scattering and ground reflection all are relatively important. Typical results for I, Q, U, and the polarization are shown in Figures 3.1.1 through 3.1.4. These results are for a receiver on the ground looking up with a zenith cosine angle of 0.4, 60° from the solar-vertical plane. As can be seen, agreement is excellent both for zero ground albedo and high ground albedo (0.8). Thus reflections at the ground are handled properly as is the basic transport through the atmosphere.

30

Å

.4 0.5 0.6 $\mu_{\rm o}$ (SOLAR ZENITH COSINE)

0.3



32

FIGURE 3.1.4.

In the next test, we add vertical structure and a polydisperse aerosol in a fairly realistic model of a clear atmosphere. Since analytical results are unavailable for such an atmosphere, we compare CTRANS' computed values with two sets of results kindly provided by Dr. R.S. Fraser (computed by programs named RADTRAN and VPD). Results are presented in Tables 3.1.1 and 3.1.2 for the intensity and polarization at the bottom and the top of the atmosphere. As may be seen, CTRANS is in agreement with both routines to within the bounds expected from CTRANS' computed standard errors.

CTRANS may encounter difficulty in handling optically thick atmospheres composed of scatterers having an extremely peaked phase function such as is provided by a Deirmandjian C1 polydispersion. The manifestation of the problem is a rather large variance especially evident when computing the radiance for a receiver looking down into the atmosphere. Tables 3.1.3 and 3.1.4 provide an illustration of the effect. Here we present a test case for a receiver pointing towards the nadir from the top of the atmosphere. The atmosphere is composed of pure Mie scatterers (C1 Deirmendjian haze: α = 6.0, β = 1.5, γ = 1.0, index of refraction = 1.5). wavelength is 0.825 microns. In the tables, $I_{\rm c}$ refers to CTRANS' results; I_k are results computed by a doubling method (provided by Mr. Lee Kyle); and SE denotes the standard error inherent in \mathbf{I}_{c} . For this case, the ground albedo was zero and polarization was neglected. As may be seen, the agreement between I_c and I_k for the τ = 0.1 case is acceptable. However, when the optical thickness is increased to $\tau = 10.0$, the variance rises markedly (though no bias is detected). This variance is extremely difficult to reduce; particularly for the τ = 10.0 case. Here 2000 photons were tracked and up to 50 scatterings allowed per track, resulting in 39,422

Table 3.1.1. Realistic Clear Atmosphere Test: Intensity

| | | | | | | | | | | | ° ≈ 20 | 2.66 mi | | | | | | | |
|----------|--------------------|-----------|-----------|----------|-----------|-----------|-----------|-----------|-----------|---|--------------------------------------|---------------------------------|---|---------------------------------|-----------------------------|---------------------|------------|-----------------------------|--|
| | IC-IVPD | +.43 | 60 | +.93 | +.19 | 65 | 70 | +.02 | 40 | iver = 10 ⁴ | oton Track | Time) = 2. | | | | | | | |
| VPD | $^{ m I}_{ m VPD}$ | .07569 | .02405 | .1361 | .02596 | .02190 | .02735 | .01455 | .02180 | ons per Rece | rings per Ph | Time (CPU | | | | | | | |
| I - I | sec | 41 | +.85 | 40 | +.44 | +.84 | +1.52 | +1.65 | +1.19 | Number of Photons per Receiver = 10^4 | Maximum Scatterings per Photon Track | Total Execution Time (CPU Time) | | | | | | | |
| RADTRAN | $^{\rm I}_{ m R}$ | .07660 | .02359 | .13947 | .02682 | .02128 | .02623 | .01396 | .02098 | Nui | Max | To | | 13 | | r<.03 | 03 < r < 2 | .2 <r< td=""><td></td></r<> | |
| SI | Std.Error (sec) | .00108 | .000488 | .002537 | .000549 | .000415 | .000504 | .000362 | .000514 | - | | . 34 | Total Rayleigh Optical Thickness $\tau_{\rm p} = 0.098$ | s τ _Λ =0.19643 | E | $1(\mathbf{r}) = 0$ |) = | $=C r^{-4}$ | |
| CTRANS | IC | .076158 + | .024003 + | .13846 + | .026063 + | .021629 + | .026991 + | .014557 + | .021593 + | m n | = 22° | Number of Atmospheric Layers = | ical Thickne | Total Aerosol Optical Thickness | Aerosol Index of Refraction | Distribution n | | | |
| | Height (Km) | 0 | 0 | 0 | 0 | 7.0 | 70 | 7.0 | 70 | Wavelength λ=0.55 μm | Zenith Angle = | Atmosphe | leigh Opt | osol Opti | I Index o | Size | | | |
| RECEIVER | Φ | 0 | 180 | 0 | 180 | 0 | 180 | 0 | 100 | lengt] | r Zen: | er of | 1 Ray | 1 Aer | eroso | Aerosol | | | |
| | Ф | 9 | 09 | 36 | 36 | 09 | 09 | 36 | 36 | Wave | Solar | Numb | Tota | Tota | A | A | | | |

Table 3.1.2. Realistic Clear Atmosphere Test: Polarization

| 1 | | | • |
|---------------------------|----------------------------|-------------------------------------|--|
| D PC-PVPD sec | 22 52 -2.51 -2.14 | -1.01 +.77 78 -1.42 | eiver=10 ⁴ hoton=20)=2.66 min. |
| VPD P _{VPD} | .0644 .5126 .0010 | .5966 .1546 .4037 0157 | ons per Rec rings per P n Time (CPU |
| PC-PR | 1.31 .39 -2.24 90 | 25 +1.04 33 90 | Number of Photons per Receiver= 10^4 Maximum Scatterings per Photon= 20 Total Execution Time (CPU)= 2.66 min. μm |
| RADTRAN I _R | .06144 .4990 .008 | .5842 .1515 .3970 | r<.03 0.03< |
| Std.Error (sec.) | .00194 .01497 .00076 | .0163 .0114 .0151 | $r_{R} = 0.098$ $r_{A} = 0.19643$ $m = 1.50$ $n(r) = 0$ $= C$ |
| CTRANS | + + + + | .+1+1+1+1 | =34 mess less on |
| CTR P.C. | .06398 .50488 000905 | .58006 .16338 .39195 02306 | Wavelength λ=0.55 μm Solar Zenith Angle=22° Number of Atmospheric Layers=34 Total Rayleigh Optical Thickness Total Aerosol Optical Thickness Aerosol Index of Refraction Aerosol Size Distribution |
| ER Height (Km) | 0 0 0 | 70 70 70 70 | Wavelength λ=0.55 μm Solar Zenith Angle=22° Number of Atmospheric Total Rayleigh Optical Total Aerosol Optical Aerosol Index of Re Aerosol Size Distri |
| RECEIVER | 0 180 0 180 | 0 180 0 180 | elength ar Zeni ber of al Rayl al Aero Aerosol |
| REC 0 | 60 60 36 36 | 60 60 36 36 | Wave Solan Numbe Tota Tota Ae |

Table 3.1.3. Uniform Aerosol Test τ =0.10

| SOLAR ZENITH | | CTRANS | KYLE | 1 | ; | PERCENT ERROR |
|-------------------------|-----------------------------|------------------------------|---|-----------|---|---|
| Angle ⁶ 0 | Intensity ^I C | Std.Error SE _C | $f{r}$ Intensity $f{I}_K$ | Sec_sec_ | $\frac{\mathrm{SE_C}}{\mathrm{I_C}} \times 100$ | $\frac{\text{IC-IK}}{\text{IK}} \times 100\%$ |
| | .00660 | + .00014 | 66900* | -2.77 | 2.1 | -5.6 |
| | .01398 | + .00024 | .01415 | 59 | 2.1 | -1.2 |
| | .00467 | + .00017 | .00472 | 32 | 3.1 | 1.1 |
| | .00620 | + .00084 | .00534 | +1.02 | 13.6 | +16.0 |
| | .000818 | + .000034 | .000812 | +.17 | | +0.7 |
| | .000368 | + .000020 | .0003726 | +.21 | | +1.1 |
| | .000360 | + .000017 | .0004017 | -2.47 | | -10.5 |
| Atmosphere: | Uniform Aerosol 5 Km Th | 5 Km Thick | | Wavelen | Wavelength λ=0.825 μ | шп |
| | Aerosol Size Di | Size Distribution: | Deirmendjian Cl | Receive | Receiver Nadir Angle=0° | 00=6 |
| | | | $\alpha = 6.0$, $\beta = 1.5$, $\gamma = 1.0$ | Receive | Receiver Altitude = | = 5 Km |
| | Aerosol Index c | of Refraction | tion m=1.5 | | | (Top of |
| | Total Vertical | Optica1 | Thickness $\tau=0.10$ | Number of | Atmo of Photon Track = | Atmosphere) ack = 22000 |
| | | | | CPU Time | CPU Time = 1.51 min. | |

Table 3.1.4. Uniform Aerosol Test r=10.0

scattering events (19.71 scatterings per photon track on the average). Tracking 2000 photons in this thick atmosphere required 5.10 minutes of CPU time. The variance might be reduced by increasing the number of photon tracks, but in general this would be very expensive. There are two effects operating in the $\tau = 10$. Case: 1) Since the atmosphere is thick, photons must scatter many times before wandering out 2) The fact that the scattering phase of the atmosphere. function is extremely peaked in the forward direction means that most photons will not contribute importantly to the sample; i.e., only a few photons which, by chance, are traveling nearly toward the sampling direction provide large contributions to the sample while all others contribute small The backwards tracking code is thus statistically inefficient in this case. This inefficiency may be unimportant in optically thin atmospheres (or in absorbing atmospheres) since it may be economically overcome by increasing the number of photon tracks.

In an attempt to improve the efficiency of the code, a number of alternate importance sampling schemes were implemented. In all of these, it was found that any substantial reduction in variance was accompanied by the introduction of bias. These alternate schemes have, thus, been abandoned.

3.2 HORIZONTAL ATMOSPHERIC INHOMOGENEITIES (CLOUDS)

In order to validate CTRANS' capacity to handle finite clouds, we may compare with the results of McKee and Cox. ³⁾ Their computation utilized a Monte Carlo method to calculate the flux emerging from the faces of a finite cloud of cubic shape. Their model ignored scattering outside the cloud and neglected polarization effects. Accordingly, we used no ambient atmosphere and restricted the ground to zero albedo.

McKee and Cox obtained results for clouds composed of an aerosol having a Deirmendjian haze Cl^{5} size distribution and optical thickness per kilometer equal to 4.9 and 73.5. The wavelength is 0.45 μm .

The scattering phase function for this case contains a very narrow, strong forward peak. Our previous considerations would lead us to expect that these clouds would provide a very severe test of CTRANS. Optically thick clouds entail many scatterings per photon track and a sharply peaked phase function requires many independent photon tracks to achieve satisfactory variance. On this account, we have restricted the test to a cloud with an optical thickness of 4.9 per kilometer.

The basic geometry of the test is illustrated in Figure 3.2.1. Light is incident upon the cloud from zenith angles of 0° , 30° , and 60° . The incident flux is normalized to 1 watt/km² normal to the beam.

In order to compare with McKee and Cox, CTRANS (using backward tracking) initiates photons from a chosen face of the cloud, distributing them uniformly over the face and distributing their initial flight directions over 2π steradians. In so doing, CTRANS might be expected to incur additional variance since some regions of the cloud will be brighter than others and some exit directions will be more important. This, then, will represent an additional source of difficulty for the backwards tracking method. For this test, separate computations are performed for each face. This means that flux will not be automatically conserved and the degree to which flux is conserved will be a test of the program.

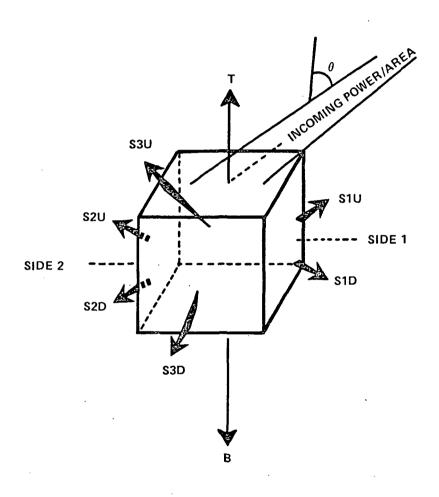


FIGURE 3.2.1. (GEOMETRY OF THE COX-McKEE TEST CLOUD

When photons are initiated on a face, they are segregated according to whether the (forward) photon would be emerging upwards or downwards. CTRANS' results for the fraction of incident flux emerging from each face upwards or downwards are compared with those of McKee and Cox in Table 3.2.1 for each face for upward-emerging and downward-emerging photons. Results for all three solar zenith angles are computed simultaneously so that there will be some correlation between results for different suns; suns at 30° and 60° correlate most strongly.

From the results of Table 3.2.1, the conservation of flux can be checked: for solar zenith angles, 0°, 30°, and 60°, the total emerging flux is $.946 \pm .12$, $.98 \pm .14$, and $1.29 \pm .25$. In each case, the difference from 1.0 is not statistically significant.

Overall, the results for individual faces tend not to be as accurate as might be desired. For a number of faces, the difference between CTRANS' results and those of McKee and Cox is large compared with the computed standard error and in most of these instances, CTRANS' results are low. may be caused indirectly by photons escaping from the cloud before they have added their full proper contributions to the flux. Two possibilities exist for correcting this: keep the photons in the cloud and accumulate the probability that they would have escaped; or run many more photons. first method precludes interaction with the atmosphere surrounding the cloud and with other clouds and is therefore generally undesirable for CTRANS. The second method will only be feasible for optically thin clouds. One other possibility exists which has not yet been fully explored: bute the photons non-uniformly over the cloud face initially.

TABLE 3.2.1. FLUX EMERGING FROM THE FACES OF A CUBIC CLOUD

| CLOUD VERTICAL OPTICAL THICKNESS = 4.9 CLOUD COMPOSITION: C1 DEIRMENDJIAN $[n\ (r) \sim r^6\exp{(-1.5\ r)}]$ | ICAL OPTICAL THICKN OSITION: C1 DEIRMEN [n (r) ~ r ⁶ exp (–1.5 r)] | KNESS = 4.9 IENDJIAN r)] | WAVI | WAVELENGTH λ = 0.45 μ m | = 0.45 µm | NUMBE | R OF PHOTON NO AMBIE GROUND | NUMBER OF PHOTONS TRACKED = 4000/FACE NO AMBIENT ATMOSPHERE GROUND ALBEDO = 0.0 | = 4000/FACE ERE 0 | |
|--|---|--------------------------------|--------------|-------------------------------------|-------------|----------------|-----------------------------------|---|-------------------------|------------|
| θ | F | S1U | S2U | รรบ | TOTAL | . 8 | S1D | .: QZS | S3D | TOTAL |
| INCIDENT | 1.00 | | | | | | | | | |
| DIRECT (CTRANS) | | | | | | .0074566 | | | | |
| (EXACT) | | | | | | .0074650 | · | | | |
| 0° SCATTERED (CTRANS) | .0561 ± .003 | .0163 ± .003 | .0165 ± .003 | .0167 ± .003 | .120 ± .006 | .2747 + .028 | .1972 + .112 | .0873 ± .012 | .1282 ± .028 | .816 ± .13 |
| (COX-McKEE) | 80: | .02 | .02 | .02 | .17 | .41 | .10 | .10 | .10 | .82 |
| INCIDENT | .63 | .37 | | | | | | | | |
| DIRECT (CTRANS) | | | | | | .06618 ± .002 | | .06633 ± .002 | | |
| 30° (EXACT) | | | | | | .06543 | | .064497 | | |
| SCATTERED (CTRANS) | .041 + .002 | .0187 ± .002 | .0162 ± .002 | .0122 + .002 | .101 ± .005 | .430 ± .135 | .0252 + .002 | .1801 ± .030 | .0560 ± .004 | .748 ± .14 |
| (COX-McKEE) | .05 | .02 | .02 | .02 | .13 | .33 | .03 | .25 | .07 | .74 |
| INCIDENT | .37 | .63 | | | | | | | | |
| DIRECT (CTRANS) | | | | | | .06494 + .0013 | | .06862 + .002 | | |
| 60° (EXACT) | | | | | | .064497 | | .06543 | | |
| SCATTERED (CTRANS) | .0428 ± .002 | .0227 ± .002 | .201 + .109 | .0211 + .003 | .309 ± .11 | .4351 ± .212 | .0234 + .008 | .2939 ± .094 | .0459 + .003 | .844 + .23 |
| (COX-McKEE) | .05 | .03 | .05 | .02 | .19 | .26 | .02 | .27 | 90: | 89. |
| | | | | _ | | | | | | |

This is tricky, however, as great care must be employed to avoid introducing bias in the final result. Further, a suitable non-uniform distribution would be expected to depend upon the solar zenith angle, thus limiting runs to a single sun. The logical extension of this concept would be to attempt to distribute photons both in position and direction according to their importance in the final results. Initial effort along this line were not notably successful and have not been pursued.

3.3 FLUX COMPUTATIONS FOR A FINITE CLOUD COMPUTED BY A FORWARD TRACKING PROGRAM

Our experience in applying the backward tracking Monte Carlo program to the evaluation of fluxes emanating through the faces of finite clouds demonstrated that there is considerable difficulty in establishing relevant importance functions which provide unbiased solutions with low variance. A more natural procedure for evaluating the fluxes through the faces of a single cloud is to use the forward tracking procedure of Cox and McKee.

In order to provide a means of testing flux results generated by the backwards code we modified an existing forward tracking code for a plane parallel atmosphere using subroutines developed for the backtracking code. The code was tested by regenerating some of the results of Cox and McKee.

The flux evaluation code, FLX, provides a capability in simulation capacity that exceeds that of the Cox and McKee code. The salient features of the simulation are as follows:

- Only one box-shaped cloud can be treated. Its
 interactions with the surrounding atmosphere,
 other clouds and the ground are neglected. The
 cloud must have its upper and lower surfaces
 parallel to the ground. Its shape need not be
 cubic.
- The cloud can have arbitrary azimuthal orientation, i.e., it is not necessary for there to be a pair of faces parallel to the solar zenith plane.
- 3. Vertical inhomogeneities within the cloud are allowed. Up to 100 vertical layers may be defined with up to 5 scattering species.
- 4. Polarization is treated using the Mueller scattering algebra.

In brief, the photon tracking and sampling algorithm is as follows:

A photon enters a face of the cloud in a random location. The face of entry is selected on the basis of the area of the face projected perpendicular to the solar direction. At each scattering we compute the face towards which the photon is travelling using the subroutine CLOUT described elsewhere in this report and evaluate the optical thickness τ_i to the face. We then sample the quantity

$$S_i = w_i I_i e^{-\tau_i}$$

where w_i = weight of the photon at the ith scattering

and I_i = intensity component of Stokes vector at ith scattering as the contribution to the appropriate flux.

The photon weight is then reduced by a fraction $(1 - e^{-\tau}i)$, i.e.,

$$w_{i+1} = w_i (1 - e^{-\tau_i})$$

and the photon is forced to make its next scattering prior to exit from the cloud. If τ_{i_s} is the actual optical distance to the next scattering then τ_{i_s} is distributed as $e^{-\tau}$ but with the constraint that $\tau_{i_s} < \tau_{i}$.

The sampled fluxes from the sides of the cloud are discriminated in terms of up-welling and down-welling photons. Obviously, only upwelling photons contribute to the flux from the upper face of the cloud and only downwelling photons contribute to the flux from the lower cloud face. By treating the point of entry to the cloud as a scattering point we are able to properly evaluate the flux contribution of attenuated but unscattered solar radiation contributing to the flux.

Each photon is tracked until its weight is reduced to an arbitrary threshold. Typically we have set the threshold at 10^{-5} . A typical simulation would involve 1000 to 10000 incident photons and standard errors in the fluxes range from 1 to 5 percent.

| | Ф | ۲ | TINF | ₽ | SIU | S2U | S3U | TOTAL UP | BINF | В | SID | S2D | S3D | TOTAL |
|-----------|----|-----|------------------------------------|--------------------------------|------------------------------------|------------------------------------|------------------------------------|--------------------------------|--------------------------------|------------------------------------|--------------------------------|------------------------------------|-----------------------|--------------------------------|
| INCIDENT | | | | 1.0 | 0 | | | | | | | | | |
| SCATTERED | 0 | 4.9 | . 222 +. 012 . 877 +. 019 | .066 +.007 .512 +.022 | . 026 +. 004 . 051 +. 009 | .026 +.004 .051 +.009 | .026 +.004 .051 | .171 +.010 .716 +.020 | .771 +.012 .118 +.016 | . 395 +. 009 . 016 +. 006 | .108 +.009 .068 +.010 | .108 +.009 .068 +.010 | .108 | .829 +.010 .287 +.018 |
| INCIDENT | | | | .63 | .36 | | | | | | | | | |
| SCATTERED | 30 | 4.9 | .292 +.013 .913 +.016 | . 006 | .031 +.005 .110 +.013 | . 020 +. 003 . 026 +. 006 | . 019 +. 003 . 045 +. 008 | .138 +.009 .618 +.021 | .705 +.013 .091 +.015 | .313 +.014 .079 | .026 +.004 .118 +.013 | . 249 +. 013 . 061 +. 010 | .0072 | .862 +.009 .378 +.019 |
| INCIDENT | | | | .36 | .63 | | | | | | | | | |
| SCATTERED | 09 | 4.9 | .470 +.015 .915 +.016 | .056 | .033 +.005 .164 +.015 | .058 +.006 .021 +.007 | . 032 +. 004 . 047 +. 008 | .212 +.011 .565 +.021 | .530 +.015 .089 +.016 | . 238 +. 013 . 106 +. 014 | .026 +.004 .191 +.016 | .264 +.013 .035 +.008 | .065 +.007 .058 | .788 +.011 .447 +.021 |

Table 3.3.1 presents a comparison of the flux fractions emanating from the various faces of the cloud. Contributions from unscattered light have been added into the diffuse contribution in computing the total downwelling flux. TINF and BINF refer to corresponding results for a cloud layer of infinite horizontal extent. Apart from these differences, Table 3.3.1 is similar in format to Table 3.2.1 and permits comparisons among the results from CTRANS, FLX, and McKee and Cox.

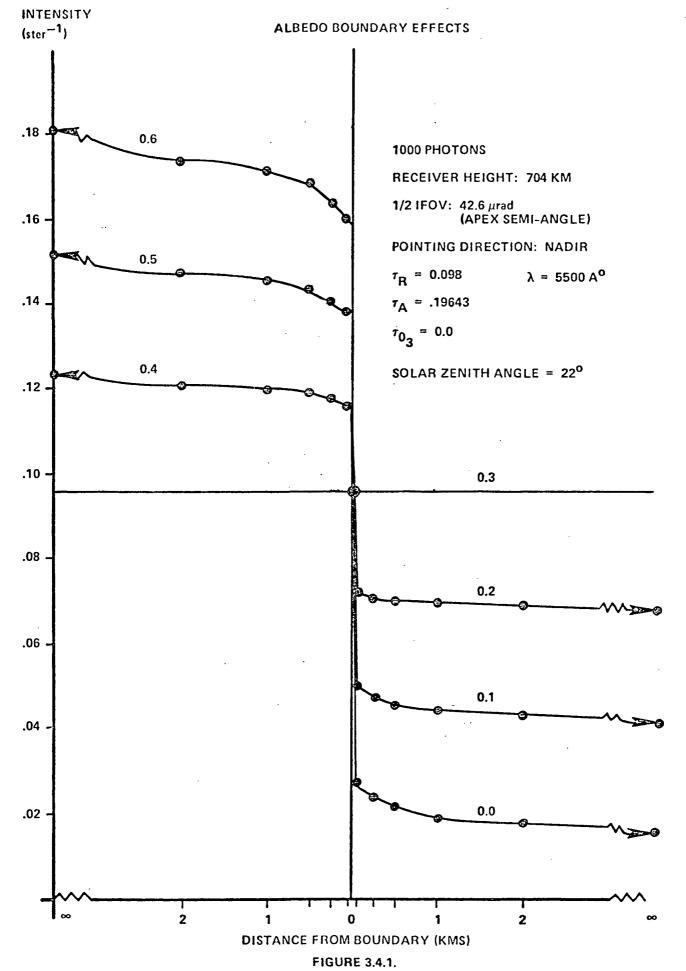
3.4 GROUND PLANE REFLECTIONS

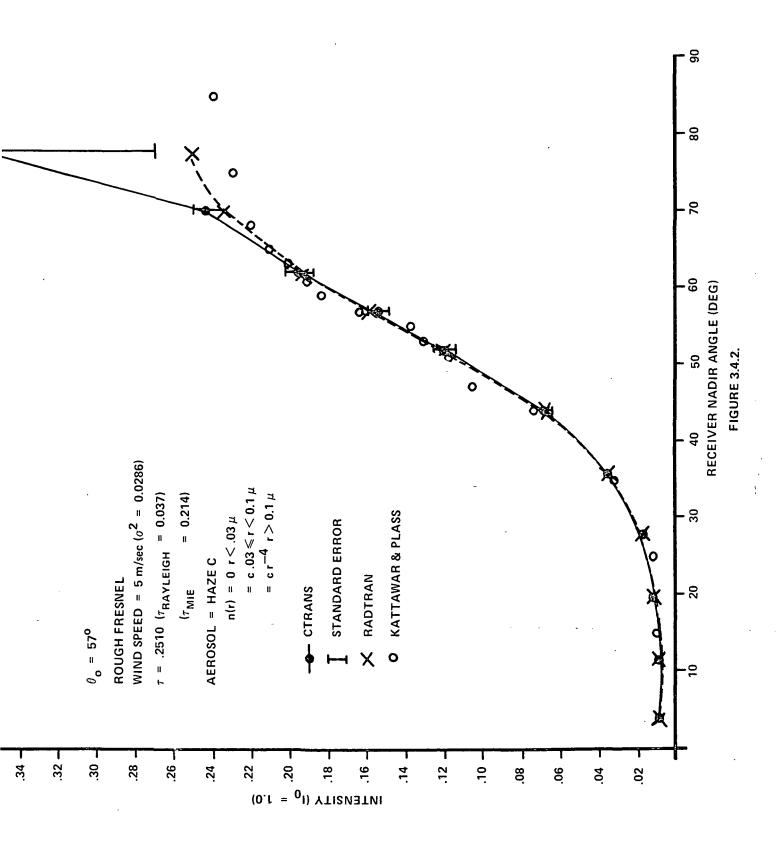
The basic mechanism for handling Lambertian reflection has been tested implicitly in the Rayleigh atmosphere tests reported in Section 3.1. The capacity of CTRANS to handle inhomogeneous Lambertian ground patterns may be illustrated by a test example but specifically comparable independent calculations do not appear to be available.

As a suitable test example, the basic ground patterns were chosen as half planes of differing albedo with boundaries parallel to the y-axis and x-axis intersections at a number of points along the axis. The receiver was above the top of the atmosphere (704 km) centered at $(X_k^R, Y_k^R) = (0,0)$ and looking towards the nadir. The receiver field of view was 85.2 µrad yielding a 60 m diameter footprint on the ground. The atmosphere was composed of Rayleigh and haze C aerosol scatterers with total optical thicknesses of .098 and .19643, respectively. The wavelength employed was $\lambda = 0.55$ µm. The albedo pairs employed were (.6, 0), (.5, .1), (.4, .2), (.3, .3) with boundaries intersecting the x-axis at $X = \pm 2.0, \pm 1.0, \pm 0.5, \pm 0.25, \pm 0.05, 0.0$ km. In addition to the half-planes, results for uniform Lambertian ground planes

with albedos .6, .5, .4, .3, .2, .1, and 0.0 were obtained All results were, of course, obtained simultaneously. These results are shown in Figure 3.4.1. Since all computed values depend upon the same set of photon tracks, relative values are computed more accurately than absolute values. This feature is of great importance when CTRANS is used, for example, to determine the distinguishability of surface features as seen from a satellite.

The ground plane reflection characteristics may be modified in yet another way: the reflectance type may be changed from Lambertian to rough Fresnel, useful for sea Several independent calculations exist surface modeling. against which CTRANS' results may be validated. We chose the following parameters to specify a suitable test case: solar zenith angle, θ_0 =57°; wind speed=5 m/sec (yields a wave slope variance $\sigma^2=0.0286$); wavelength, $\lambda=0.70 \, \mu \text{m}$; Rayleigh optical thickness τ_R =.037; aerosol optical thickness τ_{M} =0.214; the field of view of the receiver was taken to be zero. These parameter values were chosen so as to be able to compare our results with those computed independently by RADTRAN. These model parameters are also similar to (but not identical with) those used by Kattawar Figure 3.4.2 presents the results of and Plass. CTRANS as well as those of RADTRAN and Kattawar and Plass. As may be seen, all three calculations are in agreement to within 1.4 times the computed standard error.





SECTION 4.0

SUBROUTINES: DESCRIPTIVE OUTLINES

In its present version, CTRANS is composed of a MAIN control routine and 34 supporting subroutines. In the following, we outline the main functions of each of these.

MAIN

MAJOR FUNCTION: Main control routine

INPUTS: 1. Initial random number

2. IND, IST, ISKY, INFLUX, WAV

OUTPUTS: 1. Contribution maps

2. Timing information, number of photons tracked

ACTIONS: 1. Set up constants, initiate random number sequence, clear storage.

- 2. Read major control card: IND,IST,...
 (STOP if IND=0)
- Call INPUT Input model specification, write out descriptive text
- 4. Loop over photon tracks
 - 4.1 Initialize photon at receiver
 - 4.2 Sample direct sun contribution (use SEESOL)
 - 4.3 Loop over allowed number of scatterings
 - 4.3.1 Determine distance to scattering, propagate photon along path (TRACK)
 - 4.3.2 Determine scattering species
 - 4.3.3 Sample hypothetical photon from each sun (SAMPLE)

- 4.3.4 Determine scattering angle and direction (ANGLEM and JIM; or GNDREF)
- 4.3.5 Update cumulative scattering matrix following a real scattering (RENEW)
- 4.4 Clump, collect statistical measures
- 4.5 Check time remaining (terminate if insufficient)
- 5. Write contribution maps (if finite receiver field of view)
- 6. Normalize samples
- 7. Output results (OUTPUT)
- 8. Output statistics (TAB)
- 9. Return to Step 2.

INPUT

MAJOR FUNCTION: Read in model specification parameters, compute and fill reference tables, write text describing model and input quantities.

INPUTS: 1. Atmosphere description cards

- 2. Sun specifications
- 3. Size distribution parameters
- 4. Phase matrix elements (cards or tape)
- 5. Atmospheric layer heights, temperatures
- 6. Atmospheric layer heights, component densities
- 7. Cloud input specifications (CLINP)
- 8. Receiver specifications
- 9. Ground reflectance type map specifications

OUTPUTS: 1. Main title

- 2. Simulation parameters
- 3. Receiver description
- 4. Solar array description
- 5. Ground description
 - 5.1 Reflectance type map
 - 5.2 Lambertian albedo maps
- 6. Atmosphere description
 - 6.1 Characteristics of each type of scatterer in entire atmosphere
 - 6.2 Phase functions
 - 6.3 Ambient atmospheric densities, cross section
 - 6.4 Average scattering angle
 - 6.5 Cloud specifications

SAMPLE

MAJOR FUNCTION: Accumulates a sample for hypothetical photons arriving from each Sun. Reads no cards. Writes no text.

- ACTIONS: 1. Set up photon co-ordinates (if in a cloud) (TRANS 1)
 - 2. Loop over Suns
 - 2.1 Compute scattering angle and/or direction for scattering into Sun
 - 2.2 Determine scattering plane, normal to scattering plane, rotation angle (FIXALL)
 - 2.3 Multiply rotation into cumulative scattering matrix (ROTATE)
 - 2.4 Obtain matrix elements for scattering other than Rayleigh (ELEMTS or GENGND)
 - 2.5 Multiply scattering matrix into cumulative scattering matrix (SCATMI, SCATRA or GENSCA)
 - 2.6 Compute optical path length to Sun
 (SEESOL) and, thence, the attenuation
 factor
 - 2.7 Fold in other applicable weights
 - 2.8 Accumulate sample into sampling matrix, VSAMP

SEESOL

MAJOR FUNCTION: Determines optical path length to the Sun.

Assumes that CLCORD and COORD contain
position and path direction cosines for
photons in a cloud or not in a cloud,
respectively.

ACTIONS: 1. Determine distance out of cloud (if in a cloud) (CLOUT)

- 2. Compute contribution of overlying ambient atmospheric layers
- Compute contribution of present ambient atmospheric layer
- 4. Compute contribution of other clouds entered along path
 - 4.1 Check cloud spheres entered along path (CLENTR)
 - 4.2 For each sphere entered, check entries in detail; get distance to entry and distance to exit for each cloud entered (SDIST,CLOUT)
- 5. Compute total optical distance to Sun (along the direction indicated by COORD and/or CLCORD)

TRACK

MAJOR FUNCTIONS:

- Determines optical path length to be traversed
- Propagate photon along its current direction until chosen optical path length is traversed
- Update co-ordinate systems at boundary (cloud) crossings
- 4. Set IOUT=1 for photons exiting top of atmosphere
- 5. Set ISCAT=1 for photons hitting earth, compute coordinates of impact

ACTIONS:

- 1. Choose optical path length to be traversed $\tau=-\ln(P)$ (P=uniformly distributed random number)
- 2. Set up boundaries
 - 2.1 Boundary of present cloud (if in a cloud)
 - 2.2 If not in a cloud, boundary of any cloud along extended path use CLENTR, TDIST-or, atmospheric boundaries if no clouds can be entered
 - 2.3 Determine type of boundary
- 3. Successively accumulate increments in distance traversed and traversed optical path length until either the chosen optical path is exhausted or until a boundary is encountered.
- 4. Boundary encounter
 - 4.1 If boundary is top of atmosphere, set IOUT=1, exit
 - 4.2 If boundary is ground, set ISCAT=1, compute coordinates of impact, CG.
 - 4.3 If boundary is a cloud boundary, update coordinates and resume tracking
- 5. Update photon position

TRANS
TRANS
TRANS
2

FUNCTIONS: These closely related and similar routines transform coordinate positions and/or photon flight direction cosines into and out of the cloud coordinate systems.

If XC_i = photon position coordinate in the cloud i=1,3

XC_i = photon flight direction cosines in
the cloud i=4,6

X_i = photon flight direction cosines in
 the fixed coordinate system i=4,6

and XCL; = C.M coordinates of the cloud in the the fixed system i=1,3

 θ = orientation angle of the cloud

$$\underline{XC} = \underline{\underline{R}}(\underline{X} - \underline{XCL}) \qquad i=1,3$$

$$XC = \underline{\underline{R}}(\underline{X})$$
 $i=4,6$

$$\underline{X} = \underline{XCL} + \hat{R} \underline{XC} \quad i=1,3$$

$$X = \tilde{R} XC$$
 i=4,6

$$R = \begin{pmatrix} \cos\theta & \sin\theta & 0 \\ -\sin\theta & \cos\theta & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

 \tilde{R} = transpose of R

CLOUT (I, NOUT, CDIST, EXT)

FUNCTION: Determines the point of exit from a cloud and the distance along the photon path to this point

PARAMETERS: I = Cloud index

NOUT = 1 if exit point is found

CDIST(I) = distance to exit point

EXT(I,J),J=1,2,3 = exit point coordinates for cloud I

METHOD: Let C_i, C_j, C_k = exit coordinates in cloud system X_i, X_j, X_k = present photon coordinates in cloud system V_i, V_j, V_k = direction cosines of photon flight path in cloud system A_i, A_j, A_k = cloud half edge lengths all where i, J, K may be a permutation of X, Y, Z.

1. $C_i = |V_i| *A_i$ $C_j = X_j + (C_i - X_i) V_j / V_i$ $C_k = X_k + (C_i - X_i) V_k / V_i$ if C_j and C_k are within the cloud boundaries then the exit point is (C_i, C_j, C_k) i if not, cycle through values of (i, j, k)

2. Compute distance to exit point

CLINP

FUNCTIONS: Reads in cloud input parameters, transforms these and computes auxilliary descriptive parameters.

ACTIONS:

- 1. Read number of clouds NCLOUD
- 2. Return if NCLOUD=0
- 3. For each cloud read
 HC,AC,BL,CC,XCL,YCL,THETAC,INC,CLDEN
 where

THETAC =Orientation angle of the cloud measured clockwise looking up (in degrees)

CLDEN = Density of scatterers of type INC (cm^{-3}) .

4. Computes sines and cosines of THETAC, cloud edge half-lengths, radii of cloud circumscribing spheres and puts all information into the descriptive vector CLOUD(10,10):

CLOUD(I,J) : I=Cloud pointer

J=1,2,3 : C.M. coordinates of the cloud J=4,5,6 : Half dimensions of each side

of the cloud

J=7,8 : cos (THETAC) , sin (THETAC)

J=9 : Radius, R, of the circumscribing

sphere

J=10 : R^{2}

CLENTR

FUNCTIONS: Checks whether the photon can penetrate any of the cloud circumscribing spheres and orders the sequence of entry.

ACTIONS:

- 1. Computes square of distance to a cloud, length between point of closest approach and present photon position, and square of distance of closest approach along the extended path.
- 2. Orders entries by entry sequence

SDIST

FUNCTIONS: Tests which plane of the cloud photon entered, calculates distance to intersection, and orders them by entry sequence

ACTIONS:

- 1. For each uneliminated cloud computes point of entry and distance to that point (PLANE)
- 2. Orders by entry sequence (length from present position to entry point.

TDIST

FUNCTION: Calculates distance to the cloud photon can enter in TRACK

ACTION: Uses PLANE to compute distance to the first cloud entry

PLANE (I, INDC, ENT, NPLANE, DIST)

FUNCTION:

Calculates the coordinates of the point of entry of a photon into a cloud and determines the distance to this point from the photon's present position

PARAMETERS:

= Cloud index

INDC = Parameter controlling the transformation into/out of cloud coordinate
system

METHOD:

Let C_i, C_j, C_k = Entry coordinates in cloud system X_i, X_j, X_k = Present photon position coordinates in cloud system V_i, V_j, V_k = Direction cosines of photon flight path in cloud system A_i, A_j, A_k = Cloud half-edge lengths all where i,j,k are a permutation

- 1. If $|X_i| \le A_i$ try another permutation 2. If $X_i < A_i$ $C_i = -A_i$
- 2. If $X_i < A_i$ $C_i = -A_i$ $X_i > A_i$ $C_i = A_i$
- 3. If $(X_i C_i)A_i > 0$ or $V_i = 0$.
- 4. $C_{j} = X_{j} + (V_{j}/V_{i}) (C_{i}-X_{i})$ $C_{k} = X_{k} + (V_{k}/V_{i}) (C_{i}-X_{i})$

- If C_j and C_k lie within the cloud boundaries, exit point is (C₁,C₂,C₃) If not, try another permutation
- 6. Calculate distance to entry point, 7. If no entry point has been found after
 - three permutations, set NPLANE=0

FIXALL (UX, UY, UZ, CX, CY, CZ, UN, VN, WN, ANGLE, A, B, C,)

FUNCTIONS: Determines normal to the scattering plane and the angle through which the Stokes vector must be rotated

PARAMETERS: $\underline{U} = UX, \dot{U}Y, UZ$ Direction cosines of post-scattering photon $\underline{C} = CX, CY, CZ$ Direction cosines of prescattering photon

 \underline{N} = UN, VN, WN Direction cosines of normal to previous scattering plane ANGLE Necessary angle of rotation M = A,B,C Direction cosines of normal

to new scattering plane

ACTIONS:

- Checks (UX,UY,UZ) and CX,CY,CZ) for parallelism if parallel, return with (A,B,C)=(UN,VN,WN),ANGLE=0.
- 2. Computes direction cosines of normal to scattering plane, (A,B,C) via NORMAL
- 3. Computes dot product $\underline{N} \cdot \underline{M}$ and thence the rotation angle magnitude
- 4. Sets ANGLE=-ANGLE IF C (MxN) < 0.

NORMAL (CX1,CY1,CZ1,CX2,CY2,CZ2,CX,CY,CZ)

FUNCTION: Computes direction cosines of unit vector normal to the plane defined by two vectors

PARAMETERS:

$$\underline{V1}$$
 = (CX1,CY1,CZ1) Direction cosines of one vector defining plane

$$\underline{V2}$$
 = (CX2,CY2,CZ2) Direction cosines of other vector defining planes

$$N = (CX, CY, CZ)$$
 Direction cosines of NORMAL

ACTION:

- 1. Computes $N = V1 \times V2$
- 2. Normalizes N to unit magnitude

i.e.
$$N = \frac{V1xV2}{|V1xV2|}$$

JIM (COUT, CIN, CP, SP, CB, SB)

FUNCTION: Computes photon direction cosines after scattering given the direction cosines prior to scattering and the polar and azimuthal scattering angles

PARAMETERS: COUT(J) J=1,3 Photon position coordinates
J=4,6 Post-scattering direction
cosines

CIN(J) J=1,3 Photon position coordinates
J=4,6 Pre-scattering direction
cosines

CP,SP Cosine and sine of azimuthal scattering angle

CB,SB Cosine and sin of polar scattering angle

ACTION: Let θ , ϕ = polar coordinates of CIN(J), J=4,6 then

 $COUT(4) = CB*CIN(4)-SB*SP*SIN(\phi)+SB*CP*CIN(6)*$ $COS(\phi)$

 $COUT(5) = CB*CIN(5)+SB*SP*COS(\phi)+SB*CP*CIN(6)*$ $SIN(\phi)$

 $COUT(6) = CB*CIN(6) - SB*CP*SIN(\theta)$

XSECTM XS2

FUNCTION:

These essentially similar routines perform table look-up and interpolation of the phase function given a value for the scattering angle, B. XSECTM operates only on \mathbf{S}_{11} and XS2 operates on general phase matrix

elements.

METHOD:

Three point Newton's method - see for example R.W. Hamming, <u>Numerical Methods for Scientists</u> and <u>Engineers</u>, McGraw-Hill, New York, 1962; p 99.

ANGLEM

FUNCTION: Given a rectangularly distributed random number, and the scattering species index, chooses a scattering angle by table look-up. Uses the scattering angle distribution function formed from \mathbf{S}_{11} for the indicated species.

PARAMETERS: Where JK points to the scattering species

Y1(JK,100) = Probability distribution function for the angles listed by X1 computed from S_{11} .

METHOD: 2 or 3-point Newton's method

ELEMTS (BETA, S, K)

FUNCTION: Computes all the scattering matrix elements

(for an atmospheric scattering) for a

specified scattering angle BETA.

PARAMETERS: BETA = Scattering angle

S(I) = Result value of Ith matrix element

 $S(1) = S_{11}$

 $S(2) = S_{12}$

 $S(3) = S_{33}$

 $S(4) = S_{34}$

 $S(5) = S_{22}$

 $S(6) = S_{44}$

ACTIONS: Use XSECTM to obtain S(1) for hypothetical

scattering (S(1)=1 for real scattering).

. Use XS2 for remaining matrix elements.

RENEW (COUT, BETA, CB)

FUNCTION: Updates the cumulative scattering matrix following a real scattering

PARAMETERS: COUT = Six-vector describing post-scattering photon

BETA = Scattering angle

CB = cos(BETA)

ACTIONS:

- 1. Determine normal to new scattering plane and the angle between it and the normal to the previous scattering plane (or reference direction) using FIXALL
 - 1.1 Update scattering plane normal
- Post-multiply a rotation into the cumulative scattering matrix
- Obtain scattering matrix elements and accumulate a scattering by post-multiplication
 - 3.1 For ground scattering obtain
 matrix elements from LAMBRT (Lambertian scattering) or from GENGND
 (rough Fresnel scattering) and use
 GENSCA to accumulate the scattering.
 - 3.2 For atmospheric scattering, obtain
 Mie or generalized scattering matrix
 elements from ELEMTS. Accumulate
 the scattering using SCATRA (Rayleigh
 scattering-matrix elements computed),
 GENSCA (general scattering), or SCATMI
 (Mie scattering).
- 4. Update CIN to the values contained in COUT.

FUNCT (A1, B1, C1, D1, IDIST, D)

FUNCTION: Computes the value of the particle diameter

probability density function for a given

particle diameter.

PARAMETERS: A1,B1,C1 = distribution function parameters

D1 = Normalization factor

IDIST Specifies type of distribution

= 1 Junge

= 2 Deirmendjian

= 3 Log-normal

= 4 Haze-C (Junge with pedestal)

D = Particle diameter

USE: FUNCT is called as a function

ACTIONS: IDIST=1 FUNCT = $D1*D^{-B1}$ D>A1

= 0. D<A1

IDIST=2 FUNCT = $D1*D^{A1}$ exp (- $B1*D^{C1}$)

IDIST=3 FUNCT = 0. $D \le C1$

 $= \frac{D1}{D-C1} \exp\left\{-\left[A1*\ln\left(\frac{D-C1}{B1-C1}\right)\right]^2\right\}$

IDIST=4 FUNCT = 0 D<C1 $= D1 C1 < D < \Delta 1$

= D1 $= D1(\frac{A1}{D})^{B1} \qquad \begin{array}{c} C1 \leq D < A1 \\ D \geq A1 \end{array}$

GENSCA (SSIN, SOUT, SK, K)
SCATMI (SSIN, SOUT, S, K)
SCATRA (SSIN, SOUT, CB, K)

FUNCTIONS: These similar routines accumulate a new scattering by post-multiplying the scattering phase matrix into the cumulative scattering matrix for real or hypothetical (sampling) scattering.

PARAMETERS: SSIN = 16-element vector containing the elements of the 4x4 cumulative scattering matrix arrayed columnwise.

SOUT = 16-element vector of elements of cumulative scattering matrix after scattering.

S = 6-element vector of matrix element values

K = 1 Compute only first column of SOUT
 (i.e., first 4 elements) for
 hypothetical (sampling) scattering.
 All elements are not needed for
 unpolarized incident light.

K = 2 Compute all elements of SOUT.

CB = Cosine of scattering angle used by SCATRA to compute scattering matrix elements.

ACTIONS: If S_0 , is the output matrix, $S_{\underline{I}}$ is the input matrix

$$\underline{\underline{S_0}} = \underline{\underline{S_1}} \times \begin{pmatrix} s_{11} & s_{12} & 0 & 0 \\ s_{12} & s_{22} & 0 & 0 \\ 0 & 0 & s_{33} & s_{34} \\ 0 & 0 & -s_{34} & s_{44} \end{pmatrix}$$

SPECIALIZATION: SCATMI Assumes that
$$S_{33} = S_{44}$$
 and $S_{11} = S_{22}$ Assumes that $S_{11} = 1$ if K=2
$$SCATRA \ Computes \ S_{11} = \frac{3}{4} \ (1 + (CB)^2) = S_{22}$$

$$S_{12} = \frac{3}{4} \ ((CB)^2 - 1)$$

$$S_{33} = \frac{3}{2} \ CB = S_{44}$$

$$S_{34} = 0.$$
 Takes
$$S_{11} = 1 \ and \ normalizes \ other$$
 matrix elements by dividing by $\frac{3}{4} (1 + (CB)^2)$

if K=2.

ROTATE (SSIN, SOUT, PHI)

FUNCTION: Folds a rotation into the cumulative scatter-

ing matrix by post-multiplication.

PARAMETERS: SSIN = 16-element vector containing column-

wise ordered elements of the cumulative

scattering matrix

SOUT = 16-element vector containing elements

of the cumulative scattering matrix

after rotation.

PHI = Rotation angle

ACTIONS: S_0 is the cumulative scattering matrix

after rotation

 $\boldsymbol{S}_{\boldsymbol{I}}$ \quad is the cumulative scattering matrix

before rotation

 ϕ = rotation angle

 $\underline{S_0} = \underline{S_I} \times \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & \cos 2\phi & \sin 2\phi & 0 \\ 0 & -\sin 2\phi & \cos 2\phi & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$

REFLOC (CG, IRCOD)

FUNCTION: Returns a value of IRCOD indicating the

reflectance type at a point indicated by CG corresponding to the map choice fixed

by INPUT.

PARAMETERS: CG(1), CG(2) = X-Y location of point of impact

of photon on the ground

IRCOD = Reflectance type code

= 0 Lambertian

= 1 Rough Fresnel

ACTIONS:

2. If map is half-planes | IRPAT | = 1 then

for IRPAT=+1 and X>XBND IRCOD=IRSPOT
for IRPAT=-1 and X<XBND IRCOD=IRSPOT
with</pre>

IRCOD=IRBGD otherwise

3. If map consists of rectangles differing from background (IRPAT=2) then

IRCOD=IRSPOT if |X-CNTRX| < EDGEX/2
and |Y-CNTRY| < EDGEY/2
IRCOD=IRBGD otherwise</pre>

LAMBRT (COUT, S, K)

FUNCTION: Constructs matrix elements appropriate to

Lambertian scattering

PARAMETERS: COUT = Descriptor vector for scattered photon

COUT(6) = zenith cosine

S = 6-vector of matrix elements

K = 1 Indicates hypothetical scattering

K = 2 Indicates real scattering

ACTIONS: S(1) = 2*COUT(6) K=2

S(1) = 4*COUT(6) K=1

S(J) = 0 J=2,6

GNDREF (COUT, CIN, XTH, XPH, IRCOD)

FUNCTION: Generates a reflection direction from the

ground

PARAMETERS: COUT = Post-scattering photon descriptor

CIN = Pre-scattering photon descriptor

XTH, XPH = Uniformly distributed random

numbers

IRGOD = Reflectance type indicator

= 0 Lambertian

= 1 Rough Fresnel

ACTIONS: 1. For Lambertian reflection

 $\phi = 2\pi * XPH$

 $Cos(\theta) = XTH$

 $COUT(4) = SIN\theta COS\phi$

 $COUT(5) = SIN\theta SIN\phi$

 $COUT(6) = COS\theta$

COUT(J) = CIN(J), J=1,3

- 2. For rough Fresnel reflection
 - 2.1 Find slope X,Y components from Gaussian distribution through GAUSS (standard deviation=WSIG, mean=0.)
 - 2.2 Compute normal to slope (direction cosines labeled WNG(I), I=1,3)
 - 2.3 Check to see if slope is visible to incident photon (visible if $G = \sum_{J=1}^{3} CIN (J)*WNG(J) < 0.$

GENGND (COUT, CIN, IRCOD, S, K)

FUNCTION: Computes Fresnel matrix elements multiplied by statistical factors

PARAMETERS: COUT, CIN = Scattered and unscattered photon descriptors, respectively

IRCOD = Not used

S = 6-vector of matrix elements

K = 1 Hypothetical (sampling) scattering

= 2 Real scattering

ACTIONS: Let $\underline{V_o}$ and $\underline{V_i}$ be the (unit) photon propagation vectors following and prior to reflection, respectively and \underline{n} be the normal to the surface facet.

1. Compute cosine of angle of incidence:

$$\cos X = \sqrt{\frac{1 - \underline{V}_0 \cdot \underline{V}_1}{2}}$$

 Compute factor to account for non-vertical incidence of photon (modification to probability of encounter for a chosen facet slope)

FACTR = $\cos X/[(\hat{n} \cdot \hat{z}) (-\hat{z} \cdot V_I)]$

3. Compute cosine of angle of refraction

$$\cos \theta_t = \sqrt{1 - \frac{1}{f^2} (1 - \cos^2 X)}$$
 f=Index of refraction

4. Compute Fresnel factors

$$R_{11} = \frac{f \cos X - \cos \theta_t}{f \cos X + \cos \theta_t} ; R_1 = \frac{\cos X - f \cos \theta_t}{\cos X + f \cos \theta_t}$$

- 6. If K=1, compute probability of encountering proper slope; FEXP
- 7. Fold all auxilliary factors into matrix elements.

TAB (VSAMP2, VSAMP3, XMC, XMX, IND, IFL)

FUNCTION:

Compute and display mean, variance, standard deviation, and standard error for each sun, each Lambertian albedo pattern, and each Stokes' parameter (I,Q,U,V) for single scattering, multiple scattering and total. Compute correlation coefficients amongst suns.

PARAMETERS: VSAMP2 = Sample matrix VSAMP2(IQUV,ISUN,IS,1,IALB) where

IQUV Points to Stokes' parameters
ISUN Designates the sun

IS Designates direct sun contribution, single scattering, and higher order scattering.

VSAMP3 = Sample matrix of squares of clumps

XMC = Number of clumps

XMX = Factor used to normalize results to
 the number of photons actually tracked
 (when this varies from NMAX).

IND = 5 or 6 display results only for zero
 albedo

ACTIONS:

- 1. Compute normalization factors $f=(solar\ const.)/4\pi)$
- Compute and write means and standard deviations

std. dev. =
$$\sqrt{\frac{XMC}{XMC-1}}$$
 [VSAMP3*f²*XMC-(MEAN)²]

80

- 3. Compute variance = $(std. dev.)^2 = \sigma^2$ and standard error = $\sqrt{\sigma^2/XMC}$
- 4. Compute correlation coefficient matrix for results for each solar direction.

OUTPUT (VSAMP, IND, FACT)

FUNCTION:

Normalize results, write out intensity or flux results labeled by Lambertian albedo pattern. If appropriate, compute effective albedo evaluated from regression on uniform albedo patterns and modulation transfer functions.

PARAMETERS: VSAMP = Sample matrix

IND = Not used

ACTIONS:

- Compute normalization factor
 f=(solar const.)*FACT/π. If receiver
 is the face of a cloud, normalize
 to unit flux incident on the cloud
 (Cox-McKee normalization).
- Write out results (normalized) for intensity or flux. For intensity, results are direct sun contribution, single scattering, higher order scattering, and total intensity. For flux, results are direct solar contribution, upward flux, downward flux, total flux.
- 3. If the Lambertian albedo patterns contain uniform planes with A=0. and at least two others (and if the reflectance type map is uniformly Lambertian) perform a regression analysis: find AL,BL in the expression

[I(A)-I(0)]/A = AL + BL*[I(A)-I(0)]

where A refers to the albedo and I(A) is the result for albedo=A. Use AREG. Also computed are the effective albedos for other albedo patterns.

4. If appropriate (IMTF=1, set by ALBEDO) compute percent contrast from pairs of results both for intensity and effective albedo.

ALBEDO (IALB, ALBPT)

FUNCTION: Write titles or compute the albedo at the point specified by CG.

ACTIONS: 1. IALB = -J Write out description (50 character maximum width) of albedo pattern J.

IALB = +J Compute albedo at the point X=CG(1), Y=CG(2) and return this value as ALBPT.

2. If this pattern is for computing MTF's, set IMTF=1.

AREG (NALB, AL, BL)

FUNCTION: Perform albedo regression analysis assuming

FTOT (NALB) is the value corresponding to a

uniform zero albedo plane.

PARAMETERS: NALB = Total number of albedo patterns

NUNIF = Number of uniform plane albedo

patterns

AL, BL = Regression coefficients

ACTIONS: 1. Use LINFIT to compute albedo regression coefficients AL,BL

Regression [I(A)-I(A=0.)]/A=AL+BL*[I(A)-I(A=0.)]

2. Compute effective albedo (AEFF(I), I=1,NALB-1):

 $AEFF(I) = \frac{I(J) - I(NALB)}{AL + BL * (I(J) - I(NALB))}$

AEFF(NALB) = 0.

LINFIT (N,Y,X,A,B)

FUNCTION:

Performs linear regression, finding the parameters A,B which best fit the form Y=A+B*X where Y(N),X(N) are vectors of length N.

METHOD:

$$B = \frac{N \sum_{i=1}^{N} X_{i}Y_{i} - \sum_{i=1}^{N} X_{i} \sum_{i=1}^{N} Y_{i}}{N \sum_{i=1}^{N} X_{i}^{2} - (\sum_{i=1}^{N} X_{i})^{2}}$$

$$A = \frac{1}{N} \left(\sum_{i=1}^{N} Y_i - B \sum_{i=1}^{N} X_i \right)$$

SECTION 5.0 INPUT PARAMETERS

The input parameters are listed below, together with a concise description of their use, meaning, and default values (if applicable). The order of appearance is that of a proper data deck.

Initial random number (may be blank, default value = 65549)

(I7)

IND, IST, ISKY, INFLUX, WAV

(4I1, F8.5)

IND = Calculation extent indicator

- = 0 Stop calculation
- = 2 Input new clouds and receivers only
- = 4 Normal start
- = 5 Omit polarization, set scattering matrix $S_{ij} = 0$ for $i, j \neq 1$
- = 6 Calculate only for completely absorbing ground
- = 7 Completely absorbing ground + no polarization

IST = New atmosphere indicator

- = 1 Read in new atmosphere
- = 0 Use one already read in

ISKY = Sun indicator

- = 0 User specified suns (maximum of 5)
- = 1 Suns at $\theta = 0^{\circ}, 30^{\circ}, 60^{\circ}; \phi = 0$
- = 2 Suns at $\theta = 0^{\circ}, 30^{\circ}, 60^{\circ}; \phi = \pi/4$
- = 3 Suns at $\theta = 0^{\circ}, 30^{\circ}, 60^{\circ}; \phi = \pi/2$

INFLUX = Indicator specifying whether flux or intensity is
 to be calculated.

- = 0 Calculate flux at receiver
- = 1 Calculate intensity at receiver

WAV = Wavelength of radiation (in micrometers).

INDIC(5), IDIST(5), NLAYER, NSOL, NMAX, NSCA, NALB, IUNIT(5), NTOT(5)
(1011, 213, 16, 712, 513)

INDIC(k) = Indicator for type of kth species.

- > 5 If species is not present
- = 0 For Rayleigh (matrix elements computed)
- = 1 For monodisperse Mie (matrix elements read from binary tape)
- = 2 For polydisperse Mie (matrix elements computed while reading from binary tape)
- = 3 For experimental data (matrix elements read from cards)
- = 4 For ozone (absorption only)
- - = 1 For Junge distribution
 - = 2 For Deirmendjian's Haze model
 - = 3 For log-normal distribution
 - = 4 For Haze C (Junge with pedestal)
- NLAYER = Number of plane parallel layers in ambient atmosphere.
- NSOL = Number of solar zenith angles to be considered.
- NMAX = Number of photons to be tracked
- NSCA = Maximum number of scatterings allowed per photon track
- NALB = Number of Lambertian albedo maps
- IUNIT(k) = Indicator of (tape) unit from which matrix elements of k^{th} species is to be read.
- NTOT(k) = Number of diameter steps to be read for the kth (polydisperse) scattering species.

INDICA(5)

(511)

SCALE(5)

(5E11.6)

- - >0. Scale cross section of ambient atmosphere component k so that the total vertical optical thickness of component k=SCALE(k).

AR, BR

(2F8.6) AR and BR specify the Rayleigh cross section $\boldsymbol{\sigma}_R \colon$

 $\sigma_R = ARx10^{-28}x\lambda^{-BR}$, where λ is the wavelength.

ADEF, EXPFAC

(2F8.5)

ADEF = Ozone cross section $\frac{\pi}{2} = \frac{19}{2} = \frac{19}{2}$

 $\sigma_{0_3} = ADEF/2.687 \times 10^{19}$

EXPFAC = Scale height adjustment parameter for
 molecular species. If H is the effective
 scale height computed from densities as
 read from cards, the adjusted scale height
 H'=H/(1-H*EXPFAC).

```
INDOZ, INDCOZ
    (2I1)
                  = Ozone input units specification
       INDOZ
                  = 0 Ozone densities cm^{-3} are read in
                  ≠ 0 Ozone partial pressures are read in
                  = 1 Ozone densities computed from partial
                      pressures are temperature corrected.
       INDCOZ
                  Inactive parameter, no function ascribed
                  for it in this version.
SOLC
    (F10.6)
       SOLC
                  = Solar constant assumed for normalization
                  = 0. (or blank) on input, a value of 1.0
                    is adopted.
                  ≠ 0. Use the specified value of SOLC.
       If ISKY=0, solar directions must be read in as
follows:
THETA(1), PHI(1)
    (F7.4, F7.2)
THETA (NSOL), PHI (NSOL)
       THETA(I) = Solar zenith angle for I^{th} sun (in degrees)
                  = Solar azimuth from I<sup>th</sup> sun (in degrees)
       PHI(I)
NK(1), STEP(1), A1(1), B1(1), C1(1), ITAPE(1)
    (I3,4F8.3,I1)
```

```
NK(K), STEP(K),...
CSECT(K), ASECT(K)
    (2E11.4)
XA, S11(1,K), S12(1,K), S22(1,K), S33(1,K), S34(1,K), S44(1,K)
    (F6.1, 6E12.5)
XA, S11(NK, K), S12(NK, K), ...
NK(KMAX), STEP(KMAX), A1(KMAX), ... (Where KMAX is the maximum
                                         total number of atmospheric
                                         species present)
                = Number of angles for which the matrix elements
   NK(K)
                  of the K<sup>th</sup> species are to be read from tape
                   or cards.
                = Step size in x (x=\pi D/\lambda) for polydispersions
   STEP(K)
                   (K<sup>th</sup> species)
                = Value of x for monodispersions (Kth species)
   Al(J), Bl(J), Cl(J) are the parameters entering into the
                          particle size distribution function
                          for the J<sup>th</sup> scatterer type.
      For IDIST(J) = 1, the diameter distribution is
          F = D1 D^{-B1}
                              D > A1
              where D1=(B1-1) A1^{(B1-1)}
      For IDIST(J)=2
          F = D1 D^{A1} \exp(-BB1*D^{C1})
              where D1 = \frac{\text{C1} \quad \text{B1}^{(\text{A1+1})/\text{C1}}}{(2^{\text{A1+1}})\Gamma(\frac{\text{A1+1}}{\text{C1}})}
                      BB1 = B1*2^{-C1}
```

[When the Kth species' matrix elements are to be read from

cards, they appear as follows:]

```
For IDIST(J)=3
          F = \frac{D1}{D-C1} \exp \{-[A1 \log (\frac{D-C1}{B1-C1})]^2\}, D>C1
             = 0. D < C1
               where D1 = A1/\sqrt{\pi}
       For IDIST(J)=4
           F = 0.
                                D<C1
                                C1 \leq D < A1
           F = D1
           F = D1 \left(\frac{A1}{D}\right)^{B1} \qquad A1 \leq D
               where D1 = \frac{B1-1}{(B1-1)(A1-C1)+A1}
     ITAPE(J) = 0 Use NTOT(J) and STEP(J) as read from cards.
                 = 1 Read NTOT(J) and STEP(J) from Mie tape.
XH(1), TEMP(1)
     (F3.0, 2X, F3.0)
XH(NLAYER),TEMP(NLAYER)
                = Layer top height of J<sup>th</sup> ambient atmospheric
                  layer (in km.)
    TEMP(J) = Temperature of J^{th} layer (in °K)
XH(1), DENS(1,1)
     (F3.0, 2X, E9.3)
XH(NLAYER), DENS(1, NLAYER)
```

```
XH(1), DENS(2,1)
XH(NLAYER), DENS(KMAX, NLAYER)
               = Layer top height of J^{th} ambient atmospheric
    XH(J)
                 layer (in km.)
    DENS(K,J) = Density of K<sup>th</sup> species in J<sup>th</sup> layer (cm<sup>-3</sup>)
NCLOUD
    (I3)
       NCLOUD = Number of clouds to be considered.
    *** The following are read in if NCLOUD>0
HC(1),AC(1),BC(1),CC(1),XCL(1),YCL(1),THETAC(1),INC(1),CLDEN(1)
    (F8.5,5F8.3,F8.5,I3,E9.3)
HC (NCLOUD), AC (NCLOUD), ...
                            = Altitude of bottom surface of I<sup>th</sup>
       HC(I)
                              cloud (in km.)
       AC(I),BC(I),CC(I) = Edge lengths of I<sup>th</sup> cloud (in km.)
                            = Angle between the x-axis of the
       THETAC(I)
                              cloud and that of the general
                              co-ordinate system (measured
                              clockwise in degrees looking up
                              from the ground).
        INC(J)
                            = Indicator of the atmospheric con-
                              stituent appropriate to Ith cloud
                              [e.g., INC(1)=3] states that cloud
                              number 1 is composed of the species
                              specified by INDIC(3)].
                            = Density of cloud-specific scatters
       CLDEN(I)
                              of the type labeled by INC(I) (in cm<sup>-3</sup>).
```

IRCLD, IREC, RECANG, HR, CZO, RPHI, RX, RY, SZO, IAPAT, NCON (212, 7E10.4, 212)

IRCLD = Receiver extent indicator

= 0 Receiver with infinitesimal area

= J>0 Receiver extends over an entire face of cloud J.

IREC Specifies which face of the cloud is to be the receiver.

= 1 + X Face of cloud IRCLD

= 2 - X Face of cloud IRCLD

= 3 + Y Face of the cloud IRCLD

= 4 - Y Face of cloud IRCLD

= 5 + Z Face of cloud IRCLD

= 6 - Z Face of Cloud IRCLD

RECANG = Apex angle of cone of receiver field of view in milliradians.

HR = Height of receiver (ignored if receiver is a cloud face)

CZO = Zenith cosine of receiver pointing direction
 (ignored if receiver is a cloud face)

RPHI = Azimuthal angle of plane containing receiver field of view and the vertical (ignored if receiver is a cloud face)

RX,RY = X,Y co-ordinates of receiver (ignored if receiver is a cloud face)

SZO = Sine of receiver zenith angle

= 0. or blank - SZO is computed from CZO

0. CZO,, computed from SZO (ignored if receiver is a cloud face)

IAPAT = Index of albedo pattern for which the intensity is to be mapped within the receiver field of view.

DEFAULT VALUE-IAPAT=1 (Max=50)

NCON = Number of THETA BINS=Number of PHI bins for mapping of response within the receiver field of view.

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```
IRPAT, IRBGD, XBND, FN, WSPD, WCO, WC1, WSIG2
    (2I2,6E11.4)
       IRPAT
                 = Specification of reflectance type map
```

= 0 Uniform plane (= Background)

=+1 Plane differs from background for X>XBND

=-1 Plane differs from background for X<XBND

= 2 Inhomogeneties are rectangles set upon uniform background plane.

= Specification of character of background IRBGD

= 0 Background is Lambertian

= 1 Background is rough Fresnel

XBND = Boundary line for half-plane pattern

= X-co-ordinate of line parallel to Y-axis

= Index of refraction for Fresnel reflectance FN DEFAULT=1.338

WSPD = Wind speed (m/sec.) for rough Fresnel

= Parameters used to construct variance, WSIG2 WCO,WC1 If both = 0 (or Blank), use default values:

> WC0 = .0015WC1=2.54E-03

WSIG2 = Variance for Gaussian slope distribution for rough Fresnel reflectance Non-zero or non-blank entry takes precedence over value computed from WCO, WC1, and WSPD.

If IRPAT=2, the following cards are required.

NSPOT

(I2)

= Number of non-overlapping rectangular inhomo-NSPOT geneities (maximum NSPOT=20)

CNTRX(1), CNTRY(1), EDGEX(1), EDGEY(1)(4E11.4)

CNTRX (NSPOT), CNTRY (NSPOT),...

CNTRX(J), CNTRY(J) = X-Y co-ordinates of rectangle J EDGEX(J), EDGEY(J) = X, Y edge lengths of rectangle J

MIE TAPE FORMATS

MONODISPERSE

where

n is computed from A1(K) and STEP(K) read in on cards:

n = INT(A1(K) (STEP(K)+.5)

REFACT = complex index of refraction

ALPHA = size parameter α = $\pi d/\lambda$, d = diameter, λ = wavelength

QSCA = scattering cross section normalized by geometric cross section

QEXT = extinction cross section normalized by geometric cross section

eg
$$\sigma_{SCA}(cm^2) = Q_{SCA} \times \pi \frac{d^2}{4} \times 10^{-8}$$

with d in micrometers

COSBAR - not used

NK(K) = number of scattering angles for species k

XA = scattering angle (degrees)

ر - ک

S33
Scattering matrix elements for Jth scattering angle kth species
S12

S34 = -S34 as used by this code

POLYDISPERSE

NT,ST [Read only if ITAPE #0; for ITAPE = 0, values are taken from card input.]

I=1,NT REFACT, ALPHA, QSCA, COSBAR, QEXT

J=1,NK(K) {XA,AA,BB,XX3,XX4,YY,XX2

NT = number of steps in size

ST = dimensionless step size = $\frac{\pi}{\lambda}$ x δd where δd is the step size in diameter (microns)

REFACT, ALPHA, QSCA, COSBAR, QEXT as described under MONODISPERSE input

XA = scattering angle (degrees)

AA,BB - not used

XX3 = S33 XX4 = -S34YY = S11

XX2 = S12

Scattering matrix elements for size parameter ALPHA, scattering angle XA

SECTION 6.0 CTRANS PROGRAM LISTING

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| ## A K K N = D = F I * D * D * 2 * E E = D 3 = F C I (K) = R F X F Y F C * C C X S F * C E X I * D S C A) | LYDI | T=NTCT(K) F(ITAPE(K).NE.0) REAU(IU)NT.ST F(ITAPE(K).NE.0) ST=ST*AV/FI STP(K)=N | 2 C C C C C C C C C C C C C C C C C C C | 4 (1 × 1) H C (| 11 11 2 1 1 1 1 1 1 1 1 1 1 1 1 1 1 1 1 | 77(K)=[C1(K)+01(K)+01(K)+01(K)+01(K)+01(K)+01(K)+010+00+00+00+00+00+00+00+00+00+00+00+00 | 11 J=1.N | 1 (C · X) = SI I (C · X) +FF + YY D (C · X · I) = SI D (C · X · I) +FF + O (C · X · I) = SI O (C · X · I) +FF + | 4 (7 · X · 1) + 0 \(4 · X · 1) + \(7 · X · 1) + \(| | ATTERER CATA FROM EXPERIMENTAL ASSOCIATION | \$11(J.K).S | X************************************* | 2CAE | N11NLE ECT(K)=ADEF FCT(K)=CSECT(K | G1(K)=CSECT(K) TINCE | 1 CF (NC ECIES (K | CALCLLATE THBAR (1) = S11(1.K) *SIN(X(K.1)) | INT(I)=Y2(I)=X(K•I)' """ SIN1(I)=Y2(I)=X(K•I)=X(NIINUE | HSUREC. PURBC. URBO. | | 2.2 | =C*I (K*I)=X(K*K1+1) (A(K*K1+1)=A(K*K1=4- | X21=X(K,K1)=X(K,K1+1) X22=X(K,K1+1)=X(K,K1) X21=X0=X=X=X=X=X=X=X=X=X=X=X=X=X=X=X=X=X=X | トム・トートルジャント・カー |
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| | (INDUT) (INDUT) (ELIMBOLT) (=1. HYPCTHETICAL SCATTERING, COMPUTE CNLY, CCLCUND, SOUT AND USE SCATTERING MATRIX NOT LIZED CY Sit) (CLOST CHING, COMPUTE ALL CF.SCOT, USE HIGH COMPUTE COMPUTE SCATTERING MATRIX AFTER (CN SSIN(10).SOUT(10) *(1.+COZ) *(1.+COZ) *(1.52-1.) *(1.52-1.) | = \$SIN(1) + \$SIN(15) + 12 = \$SIN(1) + \$SIN(15) + 12 = \$SIN(15) + 13 = \$SIN(15) + 13 = \$SIN(15) + 13 = \$SIN(15) + 13 | = SSIN(I)*TI+SSIN(IS)*TZ = SSIN(I)*TI+SSIN(IS)*TZ A B C C C C C C C C C C C C C C C C C C | # C C S C C C S C C C S C C C S C C C S S C C C C S S C C C C S S C C C S S C C C C S S C C C S S C C C S S C C C C C S S C C C C C S S C C C C S S C C C C C S S C C C C S S C C C C C S S C C C C C S S C C C C C S S C C C C C C C C C S S C | O R I M A N C R C S S R E F E R E N C E L I S I I N G * * * * * * P C C R A N C S C A T R A C C C C P C C R A C C C C A C C C A C C C C A C C C C | E4 113 SF 1*4 6C0C B4 SGUT S XR R*4 000000 SSIN F XR R*4 0000 LABEL ADDR 20 0001AL 20 0001AL MAIN, UPT=L2, LINECNT=E2, SIZE=0 00 CK. CL.EUCDIC, NULLST, NUDECK, LC 40, WAP, NCEDIT, IC, XREF WENTS = 26, PRCGRAM SIZE = 580 |
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